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High-redshift post-reionization cosmology with 21cm intensity mapping

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Abstract. We investigate the possibility of performing cosmological studies in the redshift range $2.5 < z < 5$ through suitable extensions of existing and upcoming radio-telescopes like CHIME, HIRAX and FAST. We use the Fisher matrix technique to forecast the bounds that those instruments can place on the growth rate, the BAO distance scale parameters, the sum of the neutrino masses and the number of relativistic degrees of freedom at decoupling, N_{eff} . We point out that quantities that depend on the amplitude of the 21cm power spectrum, like $f\sigma_8$, are completely degenerate with Ω_{HI} and b_{HI} , and propose several strategies to independently constrain them through cross-correlations with other probes. Assuming 5% priors on Ω_{HI} and b_{HI} , $k_{\text{max}} = 0.2 \text{ hMpc}^{-1}$ and the primary beam wedge, we find that a HIRAX extension can constrain, within bins of $\Delta z = 0.1$: 1) the value of $f\sigma_8$ at $\simeq 4\%$, 2) the value of D_A and H at $\simeq 1\%$. In combination with data from Euclid-like galaxy surveys and CMB S4, the sum of the neutrino masses can be constrained with an error equal to 23 meV (1σ), while N_{eff} can be constrained within 0.02 (1σ). We derive similar constraints for the extensions of the other instruments. We study in detail the dependence of our results on the instrument, amplitude of the HI bias, the foreground wedge coverage, the nonlinear scale used in the analysis, uncertainties in the theoretical modeling and the priors on b_{HI} and Ω_{HI} . We conclude that 21cm intensity mapping surveys operating in this redshift range can provide extremely competitive constraints on key cosmological parameters.

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1 Introduction

The Λ CDM cosmological model describes how the small, quantum fluctuations in the primordial density field were amplified by gravity to give rise to the distribution of matter we observe today: galaxies concentrated in massive galaxy clusters, fed by thin filaments surrounded by huge empty voids. It is the most accepted framework that is able to explain a very large variety of cosmological observations with just a handful of parameters. Those parameters represent fundamental physical quantities such as the geometry of the Universe, the amount of dark energy, the sum of the neutrino masses or the properties of the Universe initial conditions.

The goal of the current and upcoming surveys is to infer the value of those cosmological parameters through cosmological observables such as the anisotropies in the cosmic microwave background (CMB), the spatial distribution of galaxies, weak lensing or the properties of the Ly α forest [1-3]. The idea behind it being that the large scale structure distribution in the Universe is a faithful tracer of the underlying dark matter density field we want to probe.

Cosmic neutral hydrogen (HI) offers another way to map the matter distribution. It can be detected in emission or in absorption (e.g. in the Ly α -forest). In emission, one of the most efficient mechanism is the production of photons due to hyperfine transition of the ground state of the hydrogen atom. In this case a photon with a wavelength of 21cm (in the rest frame of the atom) is emitted, and it can be detected on Earth with radio-telescopes.

Intensity mapping (IM) [4–12] is a technique that consists in performing a low angular resolution survey to map the 21cm emission from unresolved galaxies or HI clouds. Intensity mapping for the Ly α emission line has already been successfully applied for large scale clustering studies at high redshift using BOSS data [13]. The advantages of this technique over traditional methods, such as galaxy redshift surveys, are many. Firstly, IM can survey efficiently very large cosmological volumes. Secondly, IM is spectroscopic in nature and thus offers high radial resolution. Thirdly, it can efficiently probe the spatial distribution of neutral hydrogen from the local Universe to the dark ages. On the other hand, its practical implementation is affected by some major challenges. The most important one is the fact that the amplitude of the galactic and extragalactic foregrounds is several orders of magnitude larger than the cosmological signal. Many intensity mapping surveys (CHIME¹ [14], HIRAX²[15], Tianlai³, SKA⁴ [16], FAST⁵ [17]) are currently being built and taking data with the goal of measuring the Baryon Acoustic Oscillations (BAO) scale between $0.8 < z < 2.5$ with unprecedented precision. On the other hand, the Ooty Wide Field Array (OWFA)⁶[18] is intended to measure the 21cm HI power spectrum at redshift 3.35 [19–22].

At the same time galaxy surveys such as DESI [23], EUCLID [24], LSST [25], WFIRST [26] will detect millions of galaxies up to $z \simeq 2$ and will allow us to tightly constrain the value of the cosmological parameters.

The redshift window $2.5 < z < 5$ remains however vastly unexplored on cosmological scales. DESI will measure the BAO in the three dimensional clustering of the Ly α -forest in the range $2 < z < 3.5$ remarkably well, with a 1-2% precision, but still far from the cosmic variance limit, and HETDEX [27] will observe Ly α emitting galaxies over 400 square degrees between $1.9 < z < 3.5$ probing both BAO and redshift space distortions (RSD). Indeed, carrying out galaxy redshift surveys at high-redshifts becomes increasingly more difficult, as galaxies become fainter and fainter. IM can survey these redshifts without any limitation, and therefore, assuming its technical difficulties can be overcome, it offers a natural solution to map the high-redshift Universe. For these and other reasons, a hypothetical 21cm survey at high-redshift has been considered by the Cosmic Vision program [28] as one of the five possible experiments for the next generation of cosmological surveys.

Although in most models the contribution of dark energy to the energy content of the Universe decreases with redshift, it is important to carry out observations at redshifts $z \geq 2$ in order to probe the consistency of the model. Alternatively, extensions of the models as for example early dark energy still provide viable solutions to the cosmological constant problem and could be exquisitely probed with IM of the 21cm line [29–31]. The nature of dark matter can be tested by using IM [32]. Another possibility is to constrain extensions of the vanilla Λ CDM cosmology, e.g. to massive neutrinos and curvature. More generally, measurements at high-redshift could increase our lever arm to constrain any deviation from the standard paradigm.

At high-redshift the modeling of RSD is also more robust, as the dark matter field is more linear on a cosmological scale. Unfortunately, a major limitation of all IM measurements is the interpretation of the redshift-space clustering. This is due to the conversion factor

¹<http://chime.phas.ubc.ca/>

²<http://www.acru.ukzn.ac.za/~hirax/>

³<http://tianlai.bao.ac.cn>

⁴<https://www.skatelescope.org/>

⁵<http://fast.bao.ac.cn/en/>

⁶<http://rac.ncra.tifr.res.in>

from the measured brightness temperature of the 21cm line to the underlying dark matter distribution, which is related to the mean distribution of neutral hydrogen in the Universe. Exploiting the full potential of 21cm IM, would therefore require external information about this astrophysical parameter.

In this work we want to study how much the cosmological constraints of a 21cm survey will be affected by uncertainties in the astrophysics of neutral hydrogen. We will discuss novel ways of tackle this problem, and then focus our attention to the performances of a hypothetical 21cm survey in the redshift range $2.5 < z < 5$. One of the main goals is to find the cosmological parameters that will mostly benefit from a combination of CMB probes, galaxy survey and 21cm observations.

This paper is organized as follows. In section 2 we present the modelling of the 21cm signal and of the noise along with the Fisher matrix formalism employed for this work. The characteristics of the 21cm IM surveys considered are outlined in section 2.2. The constraints on the growth rate are shown in section 3.1. In sections 4.1 and 4.2 we discuss the constraints on two possible extensions of the baseline cosmological model: neutrino masses and effective number of neutrinos, respectively. We finally outline the main conclusions of this work in section 5.

2 Method

Estimates of the performance of a given survey in constraining cosmological parameters are usually carried out using the Fisher Matrix formalism [33]. Its main ingredients are a model for the signal and the noise one would like to measure. In our case we need to define the 21cm power spectrum, $P_{21}(k, z)$, and the noise level of a given configuration of radio antennas.

The fiducial cosmology we use is the Planck 2015 [1] best-fit Λ CDM:

$$\Omega_M = 0.3075, h = 0.6774, \Omega_K = 0, n_s = 0.9667, A_s = 2.142 \times 10^{-9}, \sigma_{8,0} = 0.828, \Omega_\nu = 0.$$

For massive neutrinos we assume a fiducial value of $\Sigma m_\nu = 0.06$ eV, and for the number of relativistic degrees of freedom $N_{\text{eff}} = 3.046$. In all the calculations with non-zero Σm_ν we keep the total matter fraction $\Omega_M = \Omega_{\text{cdm}} + \Omega_b + \Omega_\nu$ fixed. We do this by including the neutrinos as the part of CDM fraction and decreasing the fiducial CDM fraction Ω_{cdm} by $\Omega_\nu = \Sigma m_\nu / (94.07 \text{ eV}) / h^2$. Throughout the whole analysis we keep the baryon fraction Ω_b fixed. Power spectra have been computed using CAMB [34].

2.1 21cm signal model

After reionization ends we can assume most of the neutral hydrogen lives in halos and galaxies, where it remains self-shielded from the background UV ionizing radiation [35]. This means that on large enough scales, a good model for the 21cm power spectrum, $P_{21}(k, z, \mu)$, is given by [36]

$$P_{21}(k, z, \mu) = \bar{T}_b^2(z) (b_{\text{HI}}(z) + f(z)\mu^2)^2 P(k, z), \quad (2.1)$$

where b_{HI} is the linear bias of the HI field, f is the growth rate we assume to follow $f = \Omega_M(z)^{0.545}$, $\mu = k_{\parallel}/k$, $P(k, z)$ is the total matter power spectrum and \bar{T}_b is the mean brightness temperature given by

$$\bar{T}_b(z) = 180 \frac{H_0(1+z)^2}{H(z)} \Omega_{\text{HI}}(z) h \text{ mK}, \quad (2.2)$$

where $\Omega_{\text{HI}}(z) = \rho_{\text{HI}}(z)/\rho_c^0$ and ρ_c^0 is the critical density of the Universe at $z = 0$. In principle, due to the non-linear effects that tend to damp the BAO peak, eq. (2.1) should contain an additional damping term: $\exp(-k_{\perp}^2 \Sigma_{\perp}^2/2 - k_{\parallel}^2 \Sigma_{\parallel}^2/2)$, where Σ_{\perp} and Σ_{\parallel} are the non-linear damping scales in the transverse and the radial direction, respectively [37]. These damping scales decrease with redshift roughly as the linear growth factor. Since we will be considering high-redshift window, the effects of non-linearities are pushed to small scales. Furthermore, a way to ameliorate this effects is to perform a BAO reconstruction method on the observed non-linear galaxy distribution [38] or the HI intensity maps at low-redshifts [39, 40]. Applying these techniques can effectively half the damping scales, pushing the effects of non-linearities to even smaller scales. In fact, we have explicitly checked that including these effects in the redshift range we consider does not make a significant difference in our results. For simplicity, we will not include the damping term in the model for the 21cm power spectrum.

One immediately sees in eq. (2.1) the main difference with respect to a cosmological analysis of a galaxy power spectrum: any constraint on the amplitude of the fluctuations, σ_8 , of RSD parameters, e.g. f , or a combination of the two, is completely degenerate with the amount of HI in the Universe. This quantity is very uncertain in the theory modeling, so any information on it must come from independent datasets. If no external priors are available then RSD measurements with the 21cm line will not be competitive with galaxy surveys. On the other hand, constraints on the BAO scales are still safe, as they are mostly independent of the broadband shape and normalization of the power spectrum.

The good news is that several measurements of the HI mean density are available nowadays [41–44]. They mostly come from the detection in quasars (QSO) spectra of Damped Ly α (DLA) systems. DLAs are objects with high column density, $N_{\text{HI}} > 10^{20.3} \text{ cm}^{-2}$, and contain more than the 90% of the total HI in the Universe. The abundance of DLAs gives a direct estimate of $\Omega_{\text{HI}}(z)$, with current errorbars around 5% at $z < 3.5$ and around 30% at $z > 3.5$ [see e.g. 41]. In both cases there is a lot of margin for improvement, as existing low redshift measurements are limited by the noise level in the spectra, whereas at high-redshift by the low numbers of QSO spectra publicly available. In the near future surveys like DESI and WEAVE [45] will collect hundreds of thousands of QSO spectra, 10 times more than current surveys, yielding much better estimates of the abundance of HI. Motivated by this, in the remaining of this work we will assume three different priors on Ω_{HI} : 2%, 5%, 10%. We take the above priors to be redshift independent, although one could also imagine a prior degrading with increasing redshift.

The other astrophysical parameter which is quite degenerate with cosmological parameters is the HI linear bias b_{HI} . In particular, any measurement that can break the $\Omega_{\text{HI}} - b_{\text{HI}}$ would be of great value for parameter estimation with the 21cm power spectrum. Naive cross-correlation with other kinds of tracers will not help, and weak lensing is not a particularly viable option since it probes the $k_{\parallel} \rightarrow 0$ limit of the density field which for 21cm surveys is the most affected by foreground contaminants [46, 47].

What we propose, based on the arguments in [48], is to measure the clustering of the same objects one uses to measure Ω_{HI} , e.g. Lyman-limit systems (LLS) and DLAs, but weighting each of them by the value of its column density. This weighted density field can then be cross-correlated with the Lyman- α forest or the 21cm field itself to partially break the $\Omega_{\text{HI}} - b_{\text{HI}}$ degeneracy. We stress that such a measurement can already be done with existing datasets, yielding the first ever measurement of the bias of the HI field. As indeed shown in [48] the

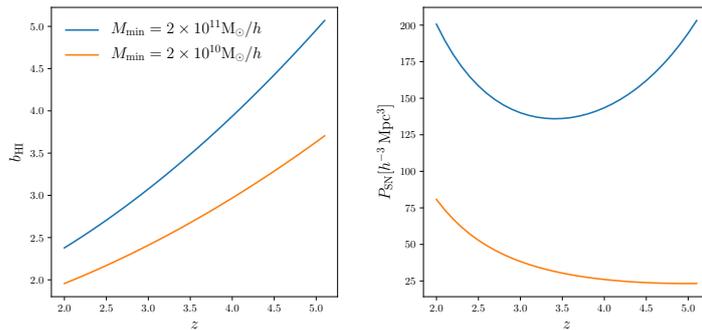


Figure 1. The figure shows the dependence on redshift of the HI bias (left) and the HI shot-noise (right) for two different models of the HI distribution (see text).

bias of the DLAs and of the HI can be written as

$$b_{\text{HI}}(z) = \frac{c}{H_0} \frac{\int_0^\infty b(M, z) n(M, z) dM \int_0^\infty d\sigma N_{\text{HI}}}{\int_0^\infty f(z, N_{\text{HI}}) N_{\text{HI}} dN_{\text{HI}}} \quad (2.3)$$

$$b_{\text{DLA}}(z) = \frac{c}{H_0} \frac{\int_0^\infty b(M, z) n(M, z) dM \int_{\sigma(10^{20.3})}^\infty d\sigma}{\int_{10^{20.3}}^\infty f(z, N_{\text{HI}}) dN_{\text{HI}}} \quad (2.4)$$

where $b(M, z)$ is the halo bias, $n(M, z)$ is the halo mass function, $d\sigma \equiv dN_{\text{HI}} \frac{d\sigma}{dN_{\text{HI}}}(N_{\text{HI}}|M, z)$ is the cross-section of systems with column density N_{HI} , i.e. the probability of observing an object with N_{HI} in a halo, and $f(z, N_{\text{HI}})$ is the column density distribution function, i.e. the abundance of such systems. We see that the difference between the two is just an extra factor N_{HI} . This means that if we weight DLAs by the value of their column density and then cross-correlate with another tracer would yield an estimate of the HI bias. The lower integration limit in eq. (2.4) is not important in this discussion, since most of the neutral hydrogen, 95% or more, lives in objects with $N_{\text{HI}} \simeq 10^{21} \text{cm}^{-2}$ [42, 48], and the weighing by N_{HI} will make the integrals above completely dominated by the tail of the column density distribution function. In the limiting case the bias of DLAs does not depend on this lower threshold, i.e. the DLA cross-section does not depend on halo mass, then $b_{\text{DLA}} = b_{\text{HI}}$. Recent measurements in [49] show a slight statistical preference for the latter situation.

The cross-correlation of the weighted DLA field, δ_{DLA}^w , with the Lyman- α forest flux, δ_F , will therefore return in the linear regime

$$\langle \delta_{\text{DLA}}^w \delta_F \rangle = b_{\text{HI}} b_F P(k) (1 + f\mu^2/b_{\text{HI}}) (1 + b_v f\mu^2/b_F) \quad (2.5)$$

with b_F and b_v the density and velocity bias of the forest. Projected correlation function will get rid of most of the RSD contribution allowing a clear determination of b_{HI} . Recent measurements of the cross-correlation between the forest and DLAs presented in [49, 50] provide an estimate of $b_{\text{DLA}} \simeq 2$ with a 5-10% error depending on different data cuts, which in the limit the DLA cross-section does not depend on halo mass exactly yields a measurement of HI bias, see eq. (2.4). Another option would be to cross-correlate δ_{DLA}^w with the 21cm field itself,

$$\langle \delta_{\text{DLA}}^w \delta_{21} \rangle = \bar{T}_b b_{\text{HI}}^2 P(k) (1 + f\mu^2/b_{\text{HI}})^2, \quad (2.6)$$

and again we see that the degeneracy between Ω_{HI} and the bias can be partially broken by combining different probes.

We expect next generation of galaxy surveys to drastically improve on these numbers. We will therefore assume, as we did above for the cosmic neutral fraction, three different redshift independent priors on b_{HI} : 2%, 5%, 10%

In a Fisher analysis we also need to specify the fiducial value of both the cosmological and astrophysical parameters. For the latter we need a model for the abundance and spatial distribution of HI, that we briefly describe here and refer the reader to [48] for further details. We assume that all HI is bound within dark matter halos and model the $M_{\text{HI}}(M, z)$ function as

$$M_{\text{HI}}(M, z) = A(z)e^{-M_{\text{min}}/M}M^\alpha, \quad (2.7)$$

where $A(z)$, M_{min} and α are free parameters. In this work we take $\alpha = 1$ and consider two values of $M_{\text{min}} = 2 \times 10^{10}h^{-1}M_\odot$ and $2 \times 10^{11}h^{-1}M_\odot$. The overall normalization, $A(z)$, is chosen such as $\Omega_{\text{HI}}(z) = \frac{1}{\rho_b} \int_0^\infty n(M, z)M_{\text{HI}}(M, z)dM = 4 \times 10^{-4}(1+z)^{0.6}$ [41], where $n(M, z)$ is the halo mass function at redshift z . The HI bias is then fully determined by the model and is given by

$$b_{\text{HI}}(z) = \frac{\int_0^\infty b(M, z)n(M, z)M_{\text{HI}}(M, z)dM}{\int_0^\infty n(M, z)M_{\text{HI}}(M, z)dM}, \quad (2.8)$$

where $b(M, z)$ represents the bias of halos of mass M at redshift z . For both the theoretical halo mass function and the halo bias we use the fitting functions calibrated from N-body simulations presented in [51]. The values of the free parameters of our model are motivated by the poorly understood redshift evolution of the HI fraction and the HI bias, both in simulations and data [35, 41, 52–55]. In Figure 1 we show the fiducial value for the bias and the effective HI shot noise $-P_{\text{SN}}$, for the two different choices of minimum mass $M_{\text{min}} = 2 \times 10^{11} M_\odot/h$ (blue) and $M_{\text{min}} = 2 \times 10^{10} M_\odot/h$ (orange). Similar halo models for the HI are also presented in [55, 56].

2.2 Noise model and 21cm IM surveys

Current and planned 21cm IM surveys focus either on the reionization epoch at $z > 6$ (e.g. HERA, PAPER, SKA) or on the low-redshift Universe $z < 2.5$, targeting the BAO feature in the 21cm power spectrum (e.g. CHIME, Tianlai, HIRAX, FAST, SKA1-MID). Some surveys (e.g. SKA1-LOW) will cover both redshift ranges but at the price of low angular resolution [57], while OWFA will focus on a single redshift [58]. Our analysis focuses on the range $2.5 < z < 5$, and therefore we first have to make a few choices on instrument specifications. We consider two approaches: interferometer instruments (similar to CHIME and HIRAX) and single-dish telescopes (similar to FAST). We modify the configuration of both types of instruments such that they can probe higher redshifts. The main requirement is to keep angular resolution good enough for the scales we want to probe, as a fixed physical distance subtend a smaller angular scale at higher redshift. This usually means we have to increase the size of the longest baseline in an interferometer or the size of the dishes in auto-correlation mode. We dub the interferometers as Ext-CHIME and Ext-HIRAX, while we dub the single-dish with highzFAST. Table 1 shows the characteristics of Ext-CHIME, Ext-HIRAX and highzFAST.

2.2.1 Thermal noise power spectrum - Interferometers

We model the thermal noise power spectrum of an interferometer as in [10],

$$P_{\text{N}}^{\text{th}}(z) = \frac{T_{\text{sys}}^2(z)X^2(z)Y(z)\lambda^4(z)S_{21}}{A_{\text{eff}}^2\text{FOV}(z)t_0n_{\text{pol}}n(\mathbf{u}, z)}, \quad (2.9)$$

	Ext-CHIME	Ext-HIRAX	highzFAST
T_a (K)	10	10	10
L_{cyl} (m)	200	-	-
W_{cyl} (m)	20	-	-
N_{cyl}	10	-	-
N_f	256	-	-
D_{dish} (m)	-	10	500
N_{dish}	-	64×64	1
η	0.7	1.0	-
D_{max} (m)	269	800	-
D_{min} (m)	L_{cyl}/N_f	10	-
N_b	2560	-	-
S_{21} (deg ²)	25000	25000	5000
n_{pol}	2	2	2
$\Delta\nu$ (kHz)	-	-	50

Table 1. Main characteristics of the considered surveys.

which we use for Ext-CHIME and Ext-HIRAX. The system temperature is the sum of antenna and the background sky temperature $T_{\text{sys}}(z) = T_a + T_{\text{sky}}(z)$. We take the antenna temperature $T_a = 10$ K and we model the sky temperature as $T_{\text{sky}}(z) = 60 \text{ K} \times (\nu_{21}(z)/300\text{MHz})^{-2.55}$ [10]. While the background sky temperature is usually below antenna temperature at low-redshifts, going to higher redshifts it increases significantly and represents the major noise contribution.

The terms $X(z)$ and $Y(z)$ are used to convert from angular and frequency space to physical space. $X(z)$ is the comoving distance and $Y(z) = \frac{c(1+z)^2}{\nu_{21}H(z)}$, where c is the speed of light, $H(z)$ is the Hubble parameter, and ν_{21} is the frequency of the 21cm line (1420 MHz). The wavelength of the 21cm line at a given redshift is $\lambda(z) = \lambda_{21}(1+z)$. We assume a survey area $S_{21} = 25000 \text{ deg}^2$, the observing time of $t_{\text{obs}} = 10000$ hours and the number of polarisation states $n_{\text{pol}} = 2$ for both Ext-CHIME and Ext-HIRAX.

The last term in the denominator is the number density of baselines in visibility space. We will assume a uniform distribution of $n(u; z) = N_b(N_b - 1)/2\pi/(u_{\text{max}}^2(z) - u_{\text{min}}^2(z))$ between $u_{\text{min}} = D_{\text{min}}/\lambda(z)$ and $u_{\text{max}} = D_{\text{max}}/\lambda(z)$. We will use an approximation of this expression $n(u; z) \simeq N_b^2/2\pi/u_{\text{max}}^2(z)$. The total number of beams N_b is the total number of antenna feeds $N_f \times N_{\text{cyl}}$ in the case of Ext-CHIME, while for Ext-HIRAX it is the number of dishes N_{dish} .

Ext-CHIME The effective area per feed is computed as $A_{\text{eff}} = \eta L_{\text{cyl}} W_{\text{cyl}}/N_f$, where η is the efficiency factor that we take $\eta = 0.7$; L_{cyl} and W_{cyl} are the length and the width of the cylinders, respectively, and N_f is the number of feeds per cylinder. While the primary beam of each element in this set-up is anisotropic, we will assume an isotropic field of view (FOV) for each element and we will use $\text{FOV} \approx 90^\circ \times 1.22\lambda(z)/W_{\text{cyl}}$ [10].

Ext-HIRAX The effective area per beam is $A_{\text{eff}} = \pi(D_{\text{dish}}/2)^2$ where D_{dish} is the diameter of a single dish. The primary beam of each dish is isotropic and defined as $\text{FOV} = \left(1.22 \frac{\lambda(z)}{D_{\text{dish}}}\right)^2$ [11].

2.2.2 Window functions for interferometers

In the case we consider – uniform number density of baselines – the noise power spectrum is scale-independent. Nevertheless, this will not be the real case and there will be a range of modes that are going to be (un)observable. To account for this we include sharp cut offs in the k -range probed by the survey, both in k_{\parallel} and k_{\perp} . We take the limits for k_{\perp} according to [10]:

$$k_{\perp}^{\min}(z) = k_{D_{\min}}(z) = \frac{2\pi D_{\min}}{D(z)\lambda(z)}, \quad (2.10)$$

$$k_{\perp}^{\max}(z) = k_{D_{\max}}(z) = \frac{2\pi D_{\max}}{D(z)\lambda(z)}, \quad (2.11)$$

where $D(z)$ is the comoving distance. The values for D_{\min} and D_{\max} depend on the configuration we consider and are given in Table 1. For $k_{\parallel}(z)$ we take the upper limit $k_{\max}(z)$ to be $0.2 h\text{Mpc}^{-1}$ or the non-linear scale $k_{\text{nl}}(z) = k_{\text{nl},0}(1+z)^{2/3}$, where $k_{\text{nl},0} = 0.2 h\text{Mpc}^{-1}$ [59]. We take the lower limit $k_{\parallel}^{\min}(z) = 2\pi/V_{\text{sur}}^{1/3}(z)$. Based on these, we can now define a window function that defines the range of scales available to a given configuration.

$$W(z, k, \mu) = \Theta(k_{D_{\max}} - k_{\perp})\Theta(k_{\perp} - k_{D_{\min}})\Theta(k_{\parallel} - k_{\parallel}^{\min})\Theta(k_{\text{nl}} - k\mu)\Theta(k_{\text{nl}} - k),$$

where Θ is the Heaviside step function. We show the effect of this window function on the signal power spectrum in Figure 2 in the case of Ext-HIRAX and in the bottom left panel of Figure 3 in the case of Ext-HIRAX and Ext-CHIME.

2.2.3 The foreground wedge

The 21cm cosmological signal is several orders of magnitude weaker than the astrophysical foreground emissions. Isolating the cosmological signal will be the major challenge in the data analysis, and it mostly relies on the assumption that foregrounds are spectrally smooth. In principle this means that only the smallest radial wave-numbers k_{\parallel} are affected by the foregrounds, allowing us to clean the data without much of the cosmological information being lost. However, the chromaticity of the interferometers response function will cause foregrounds to leak into the high- k_{\perp} modes. This will make the foreground cleaning very difficult. This effect is limited to a certain region in $k_{\parallel} - k_{\perp}$ space, acting on low- k_{\parallel} and high- k_{\perp} , and it is known in the literature as the foreground *wedge* [39, 60–65]:

$$k_{\parallel} < \sin(\theta_{\text{FoV}}) \frac{D(z)H(z)}{c(1+z)} k_{\perp}, \quad (2.12)$$

where θ_{FoV} is the field-of-view (or the primary beam width) of an interferometer element. We can also express the wedge region in terms of all the modes having $\mu < \mu_{\min}$ with:

$$\mu_{\min}(z) = \frac{k_{\parallel}}{\sqrt{k_{\parallel}^2 + k_{\perp}^2}} = \frac{\sin(\theta_{\text{FoV}})D(z)H(z)/(c(1+z))}{\sqrt{1 + [\sin(\theta_{\text{FoV}})D(z)H(z)/(c(1+z))]^2}}. \quad (2.13)$$

It is important to stress that the foreground wedge is not an intrinsic limitation of a 21cm interferometric survey, but rather a manifestation of our incomplete knowledge of the calibration of the antennas. Accurate calibration is a difficult task of any data analysis, in IM in particular because of the 10^5 difference in signal strength between the foregrounds and the signal one is trying to measure.

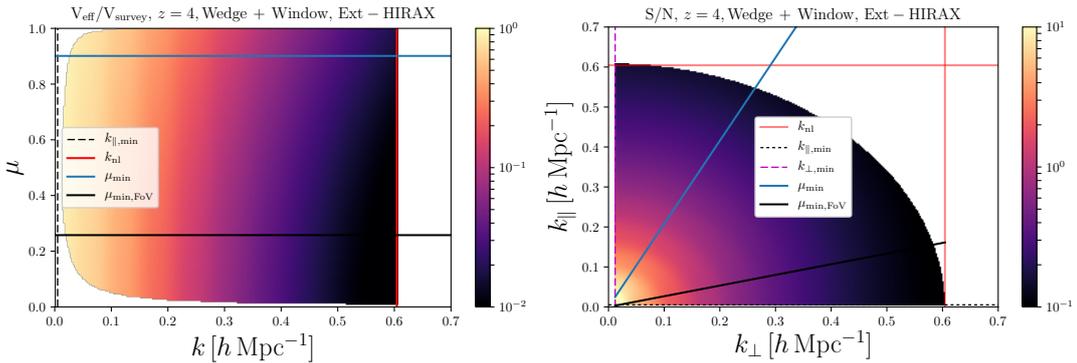


Figure 2. This plot shows the effect of the window function, wedge and k_{nl} on the number of modes available to extract cosmological information for Ext-HIRAX. The left panel displays the ratio between the effective volume to the survey volume as a function of k and μ . The right panel represents the signal-to-noise ratio as a function of k_{\perp} and k_{\parallel} . For both panels we have considered a redshift bin $\Delta z = 0.1$ centered at $z = 4$ and assumed $k_{\text{max}} = k_{\text{nl}}(z)$. Unavailable modes are shown as white regions. The effect of the wedge can be seen from the solid black and blue lines. If no wedge is present all modes shown in the panels contribute, while only the modes above the black/blue line contribute in the case of full/mid wedge.

Unfortunately, this contaminated region grows at higher redshifts, as both the comoving distance and the Hubble parameter increase faster than $(1+z)$. For instance, at $z = 2$ the foreground wedge contaminates all wave modes with $\mu \leq \mu_{\text{min}}(2) = 0.77$, and the situation gets much worse at $z = 4$ with $\mu \leq 0.9$. The simplest approach is to ignore the modes inside the wedge in data analysis (or in Fisher forecasts). This would mean that we discard 90% of the available modes at $z = 4$. Another possibility is to develop techniques that can recover as much as possible information inside the wedge [46, 47]. While these methods strongly depend on the details of the instrumental set up and of the data analysis procedure, in this paper, for simplicity we will consider few idealised cases:

1. Wedge – we discard all the modes inside the wedge assuming a horizon limit, i.e. $\sin(\theta_{\text{FoV}}) = 1$;
2. Mid-wedge – (i) in the case of Ext-HIRAX we use $\mu = \mu_{\text{min}}$; (ii) in the case of Ext-CHIME where the primary beam is anisotropic we assume modes with $\mu = \mu_{\text{min}}/1.5$ are unavailable, i.e. one third of the contaminated region is cleaned;
3. No wedge case – the foreground cleaning works perfectly for all the modes.

In Figure 2 we show the effect of the wedge and the window function on the available modes in the case of Ext-HIRAX. In the left panel we show the effective to survey volume ratio as a function of k and μ . Available modes in the three cases we consider are above the solid lines of constant μ_{min} . We show the μ_{min} in the wedge and mid-wedge case in solid blue and solid line, respectively. Similarly, in the right panel, we show the signal-to-noise ratio as a function of k_{\perp} and k_{\parallel} , where the effect of the wedge is more clearly seen. Again, solid blue and black line represent the wedge cases we consider and the available modes are on restricted on the left side of these lines. In both of the panels we use $k_{\text{max}} = k_{\text{nl}}$.

2.2.4 Thermal noise power spectrum - highzFAST

We model the noise power spectrum of a single dish as [66]

$$P_{\text{N}}^{\text{th}}(k, \mu, z) = \sigma_{\text{pix}}^2(z) V_{\text{pix}}(z) W^{-2}(k_{\perp}, z). \quad (2.14)$$

The first term is the pixel thermal noise given by

$$\sigma_{\text{pix}}^2 = \frac{T_{\text{sys}}^2}{\Delta\nu t_{\text{obs}} (\Omega_{\text{pix}}/S_{21}) N_{\text{dish}} N_{\text{beam}}}, \quad (2.15)$$

where T_{sys} is the system temperature described above, $\Delta\nu$ is the frequency resolution, t_{obs} is the total observing time, Ω_{pix} is the pixel area, S_{21} is the survey sky area, while N_{dish} and N_{beam} are the number of dishes and number of beams per dish, respectively. The pixel area is calculated assuming a Gaussian beam $\Omega_{\text{pix}} = 1.13\theta_{\text{FWHM}}^2$, where $\theta_{\text{FWHM}} = \lambda(z)/D_{\text{dish}}$. The volume of a pixel V_{pix} is computed by integrating the comoving volume element of the pixel area along the line-of-sight in the redshift range corresponding to the frequency resolution $\Delta\nu$:

$$V_{\text{pix}} = \Omega_{\text{pix}} \int_{z-\Delta z/2}^{z+\Delta z/2} dz' \frac{cD^2(z')}{H(z')}. \quad (2.16)$$

The last term in the noise power spectrum is the response function describing the angular resolution of a single-dish telescope. Since the frequency resolution we assume here is very high, we will not take into account the radial response and will only consider the angular one. We use

$$W^2(k_{\perp}, z) = \exp \left[-k_{\perp}^2 D_c^2(z) \left(\frac{\theta_{\text{FWHM}}}{\sqrt{8 \ln 2}} \right)^2 \right]. \quad (2.17)$$

For an instrument like the highzFAST the scales which will be accessible are

$$k_{\perp}^{\text{min}}(z) = \frac{2\pi}{\sqrt{D(z)^2 S_{21}}}, \quad (2.18)$$

$$k_{\perp}^{\text{max}}(z) = \frac{2\pi D_{\text{dish}}}{D(z)\lambda(z)}. \quad (2.19)$$

We show the noise power spectrum of highzFAST at $z = 4$ in the bottom right panel of Figure 3 in dashed line. The dashed gray line is the signal power spectrum, while the blue solid line includes the effect of the accessible modes.

2.3 Total noise power spectrum

The total noise power spectrum is therefore equal to

$$P_{\text{N}}^{\text{tot}} = P_{\text{N}}^{\text{th}}(z) + P_{\text{SN}}(z) \bar{T}_b^2(z). \quad (2.20)$$

We plot the different noise contributions for a survey like Ext-HIRAX at $z = 4$ in the top-left panel of Figure 3. In the top-right panel we show the expected signal-to-noise ratio for different μ values and different values of M_{min} in the case of Ext-HIRAX. Bottom panels show the total noise power spectra for both interferometers (left panel) and single-dish (right panel) in dashed lines at $z = 4$ and fixed value of M_{min} . As expected the thermal noise is the major source of uncertainties, especially going to high-redshift. We see that on cosmological

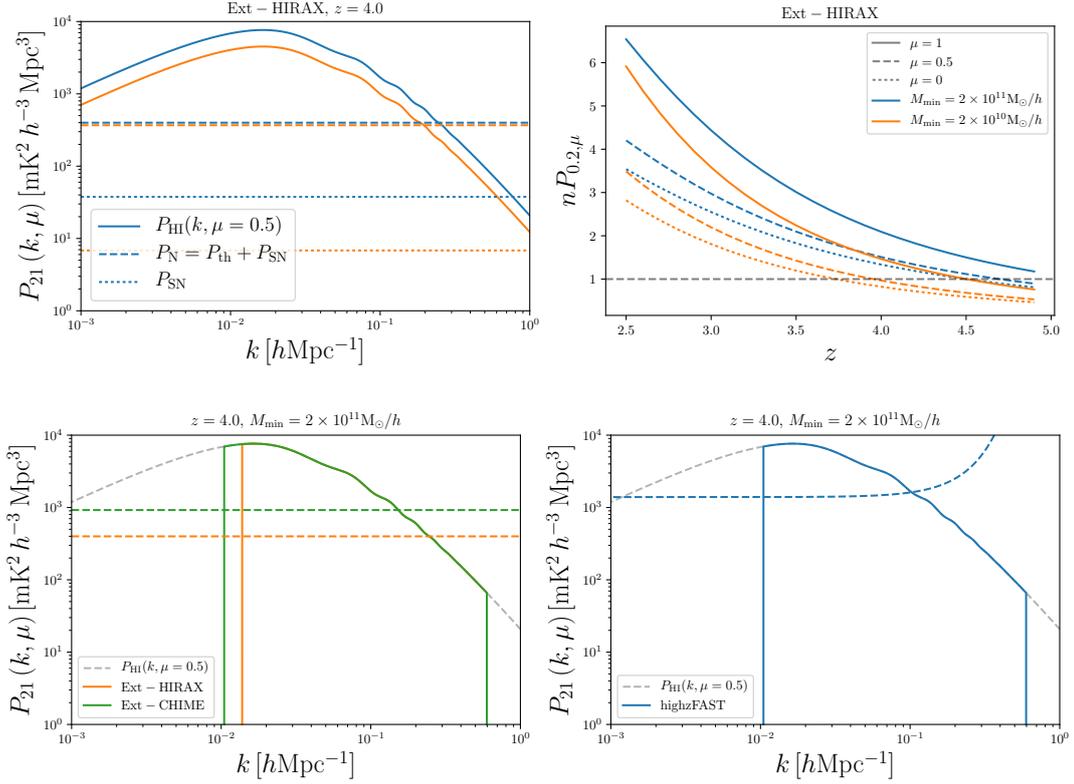


Figure 3. *Top-left:* HI power spectrum at $\mu = 0.5$ (solid lines) together with the total noise (dashed) and shot-noise (dotted) for Ext-HIRAX at $z = 4$. The different colors show the results for different values of $M_{\min} = 2 \times 10^{10} h^{-1} M_{\odot}$ (orange) and $2 \times 10^{11} h^{-1} M_{\odot}$ (blue). *Top-right:* Signal-to-noise ratio at $k = 0.2 h \text{Mpc}^{-1}$ for different values of μ , 1 (solid), 0.5 (dashed) and 0 (dotted), for the two models with different values of M_{\min} . *Bottom-left:* Effects of the window function on the HI power spectrum for $\mu = 0.5$ for Ext-HIRAX (orange) and Ext-CHIME (green) at $z = 4$. The dashed lines show the amplitude of the noise power spectrum. *Bottom-right:* Same as before but for highzFAST.

scales smaller than $k \simeq 0.2 h \text{Mpc}^{-1}$ all surveys we consider become noise dominated. In the upper left panel we plot the signal-to-noise at $k \simeq 0.2 h \text{Mpc}^{-1}$ for different values of μ

$$nP_{0.2}(z) \equiv \frac{P_{21}(k, \mu; z)}{P_{\text{N}}^{\text{tot}}}. \quad (2.21)$$

Since the shot-noise is always much smaller than the thermal noise in the antennas, a higher M_{\min} , i.e. a higher HI bias, results in a higher signal to noise ratio.

2.4 Fisher Matrix formalism

Given the signal and noise models discussed in the previous section we can use the Fisher matrix formalism to forecast the constraints on a set of parameters of interest $\{p\}$ from a 21cm IM survey. The Fisher matrix for a single redshift bin is given by [33]

$$F_{ij} = \frac{1}{8\pi^2} \int_{-1}^1 d\mu \int k^2 dk \frac{\partial \ln P_{21}(k, \mu)}{\partial p_i} \frac{\partial \ln P_{21}(k, \mu)}{\partial p_j} V_{\text{eff}}(k, \mu). \quad (2.22)$$

The effective volume is related to the comoving survey volume in a redshift bin V_{sur} by

$$V_{\text{eff}}(k, \mu) = V_{\text{sur}} \left(\frac{P_{21}(k, \mu)W(k, \mu)}{P_{21}(k, \mu)W(k, \mu) + P_{\text{N}}^{\text{tot}}(k, \mu)} \right)^2. \quad (2.23)$$

The instrument response/window function $W(k, \mu)$ is included to account for the modes that are unavailable to a given instrument. The total noise $P_{\text{N}}^{\text{tot}}$ accounts for the instrument thermal noise and the shot-noise of the HI sources as defined in eq. (2.20).

The total redshift window of the 21cm IM surveys considered here, $2.5 < z < 5$, is divided into smaller redshift bins, in each of them we calculate the Fisher matrix for a set of parameters. Under the assumption that each redshift bin is independent, the sum of the Fisher matrices in each redshift bin gives the total Fisher matrix.

Having obtained the Fisher matrix for a given redshift bin or the total redshift window, the forecasted constraint on a given parameter p_i marginalised over all the other parameters is related to the Fisher matrix by $\sigma(p_i) = \sqrt{[F^{-1}]_{ii}}$. We quote these numbers as the 1σ constraints throughout the paper.

In observations one needs to convert angular coordinates and redshifts to wave vectors in Fourier space. That is done by assuming a fiducial cosmology, which could be different from the true underlying one. By using a cosmology different from the fiducial one, we introduce an additional artificial anisotropy in the clustering measurements. This effect, known as Alcock-Paczynski(AP) effect [67], must be taken into account in the forecast. We include all the geometrical distortions arising from the assumption of a fiducial cosmology [68, 69]. The fiducial wavenumber vector \vec{k} is determined by two numbers, namely, the amplitude k_f and by the cosine of the angle between the \vec{k} and the line-of-sight, $\mu_f = k_{\parallel, f}/k_f$. The component along and transverse to the line-of-sight are related to the true values by $k_{\parallel} = k_{\parallel, f}(H/H_f)$ and $k_{\perp} = k_{\perp, f}(D_{A, f}/D_A)$. Furthermore the amplitudes of the wavenumber vectors \vec{k} and \vec{k}_f are related

$$k^2 = \left((1 - \mu_f^2) \frac{D_{A, f}^2}{D_A^2} + \mu_f^2 \frac{H^2}{H_f^2} \right) k_f^2. \quad (2.24)$$

We write our full signal power spectrum in a redshift bin centred at redshift z as [70]:

$$P_{21}(k_f, \mu_f, z) = \bar{T}_b^2(z) \frac{D_A(z)_f^2 H(z)}{D_A(z)^2 H(z)_f} (b_{\text{HI}}(z) + f(k, z)\mu^2)^2 P(k, z). \quad (2.25)$$

We will also use another way of writing the full power spectrum, particularly in the case we are interested in the redshift-space distortions where an important parameter is $f\sigma_8$:

$$P_{21}(k_f, \mu_f, z) = \bar{T}_b^2(z) \frac{D_A(z)_f^2 H(z)}{D_A(z)^2 H(z)_f} (b_{\text{HI}}\sigma_8(z) + f\sigma_8(z)\mu^2)^2 \frac{P(k, z)}{\sigma_{8, f}^2}. \quad (2.26)$$

Within our model, the derivative of the HI power spectrum with respect to a generic parameter X looks like

$$\begin{aligned} \frac{\partial \ln P_{21}}{\partial X} &= \frac{\partial \ln \bar{T}_b^2}{\partial X} + \frac{\partial \ln D_A^{-2}}{\partial X} + \frac{\partial \ln H}{\partial X} + \frac{\partial \ln P(k)}{\partial X} + \frac{\partial \ln P(k)}{\partial k} \frac{\partial k}{\partial X} \\ &+ \frac{2\mu^2}{b_{\text{HI}} + f(k)\mu^2} \left[\frac{\partial f(k)}{\partial X} + \frac{\partial f(k)}{\partial k} \frac{\partial k}{\partial X} + 2f(k)(1 - \mu^2) \left(\frac{\partial \ln H}{\partial X} + \frac{\partial \ln D_A}{\partial X} \right) \right], \end{aligned} \quad (2.27)$$

where

$$\frac{\partial k}{\partial X} = k \left[\mu^2 \frac{\partial \ln H}{\partial X} - (1 - \mu^2) \frac{\partial \ln D_A}{\partial X} \right]. \quad (2.28)$$

The Fisher matrix formalism can also accommodate theoretical errors in the signal model, for instance due to poor knowledge of non-linearities, along the way described in [71, 72]. We do not include such terms in our analysis, but in Section 4.1.1 we will comment on the effect of marginalising over uncertainties in non-linear models using templates.

3 Results from IM alone

In this section we present the constraints on the value of the cosmological parameters that can be achieved by using 21cm data alone. We focus our analysis on the growth rate and on the BAO distance scale parameters.

3.1 Growth of structures

We will first consider the case in which we assume that the geometry is fixed, e.g. by CMB measurements, and we can fit for the amplitude of the power spectrum using two parameters only: $b\sigma_8(z)$ and $f\sigma_8(z)$. This choice is motivated by a comparison with the forecasts presented in [23, 73, 74], where BAO distance scale parameters have been held fixed when quoting constraints on RSD. Following the discussion in Section 2.1, such a measurement in an IM survey is possible only with external information on Ω_{HI} and b_{HI} , that we therefore include as priors in our Fisher forecasts. We present results for 2% 5% and 10% priors in each redshift bin, but it should be kept in mind that, in reality, these priors will be a function of redshift, worsening at higher z . We show the outcome of the Fisher calculation for both a conservative choice of $k_{\text{max}} = 0.2 h \text{ Mpc}^{-1}$ and for $k_{\text{max}} = k_{\text{nl}}$.

Results are summarized in Figure 4. Starting from the continuous lines, which show the no-wedge case for Ext-CHIME, Ext-HIRAX, and highzFAST, we see that in the absence of foreground contamination 21cm surveys can measure RSD parameters to exquisite precision. Interferometric experiments saturate the priors up to $z = 3.5$, and are still able to measure $f\sigma_8$ with 4% precision at $z = 5$ in the case of aggressive priors.. In our analysis Ext-HIRAX performs better than the other two hypothetical facilities, mostly because of a better angular resolution. Single dish experiments like highzFAST pay the price of low angular resolution, but they are unaffected by the wedge, so the blue line in Figure 4 can be considered the final result for auto-correlation experiments. If we now turn our attention to the case with partial wedge contamination, the mid-wedge case described in Section 2.2.3, we see a degradation in the constraining power, with surveys with larger FOV yielding larger error bars as expected. For instance we find similar results comparing Ext-CHIME and highzFast. If the wedge extends up to the horizon the picture changes dramatically and interferometer performs much worse than single dishes. As it can be noticed in the right and lower panels, the inclusion of all the modes down to the non-linear scale k_{nl} does not substantially change the results. This is a consequence of the very low S/N ratio on scales smaller than $k \simeq 0.2 h \text{ Mpc}^{-1}$ (see Figure 3). Only in the case where we consider the full wedge, we find marginal improvement from the inclusion of smaller scales.

As discussed in Sec. 2.1, the value of the HI bias as a function of redshift is quite uncertain, reflecting our poor knowledge of what halos host neutral hydrogen. Figure 5 shows the dependence of our results on the HI bias. We do that by assuming two different values for M_{min} : $M_{\text{min}} = 2 \times 10^{10} M_{\odot}/h$ (left panel) and $M_{\text{min}} = 2 \times 10^{11} M_{\odot}/h$ (right panel).

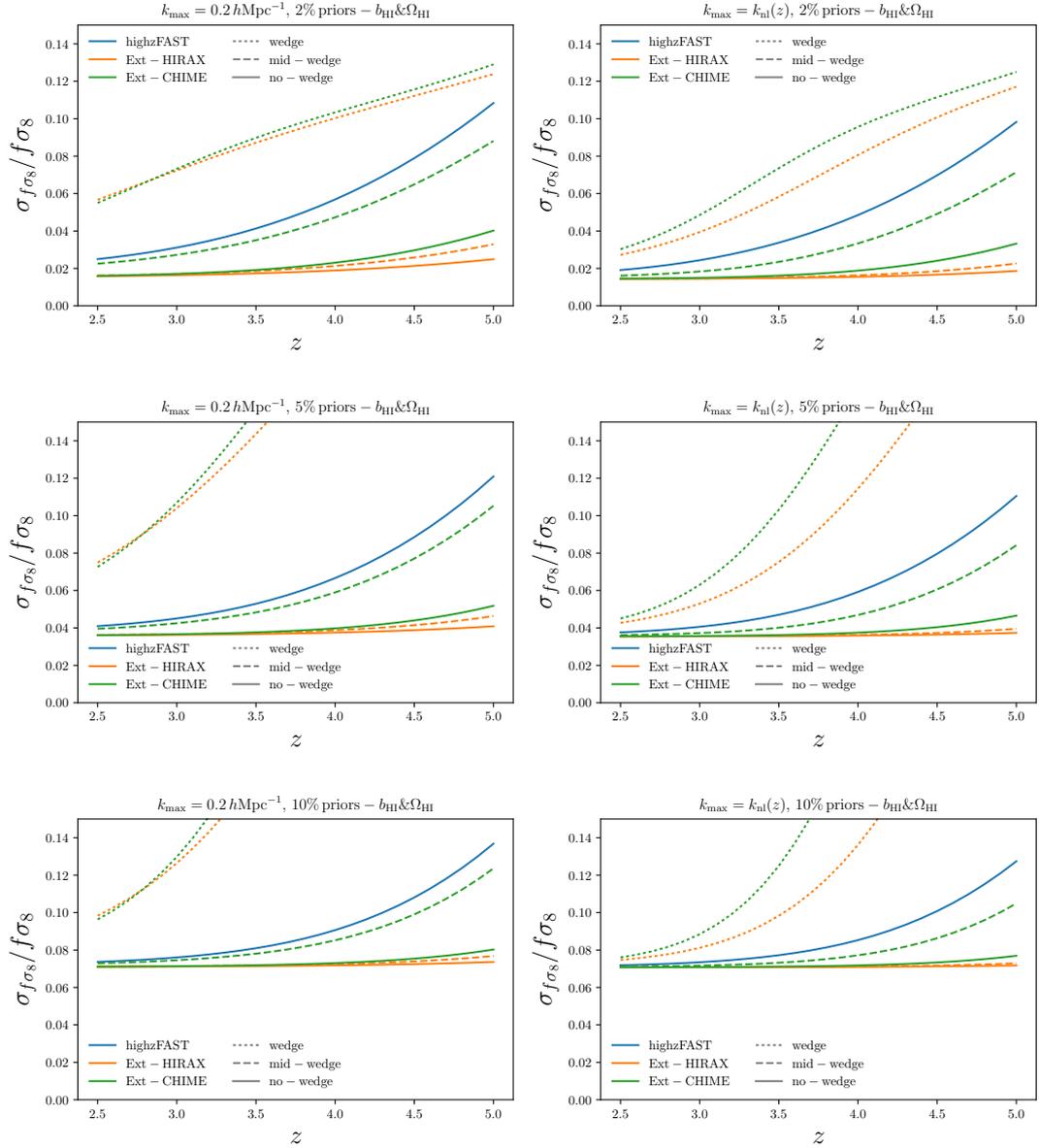


Figure 4. 1σ constraints on $f\sigma_8$ for highFAST (blue), Ext-HIRAX (orange) and Ext-CHIME (green) for different wedge configurations: no-wedge (solid), mid-wedge (dashed) and full wedge (dotted). Notice that the wedge only apply to interferometers. The different panels show the results for different assumptions of k_{\max} , $0.2 h\text{Mpc}^{-1}$ (left column) and k_{nl} (right column), and for different priors on both b_{HI} and Ω_{HI} , 2% (top row), 5% (middle row) and 10% (bottom row).

The numbers we obtain are very similar in most cases, with the exception of the full wedge scenario, indicating our forecasts should be robust to different choices of fiducial astrophysical parameters. For completeness, the numbers with the errors on $f\sigma_8$ for Ext-HIRAX and various considerations using 5% priors on the values of both Ω_{HI} and b_{HI} are shown in Table 2.

To better illustrate the synergies of a 21cm IM probe with galaxy surveys, in Figure 6 we plot, as a function of redshift, the forecasted constraints on $f\sigma_8$ from: DESI (Tables 2.3

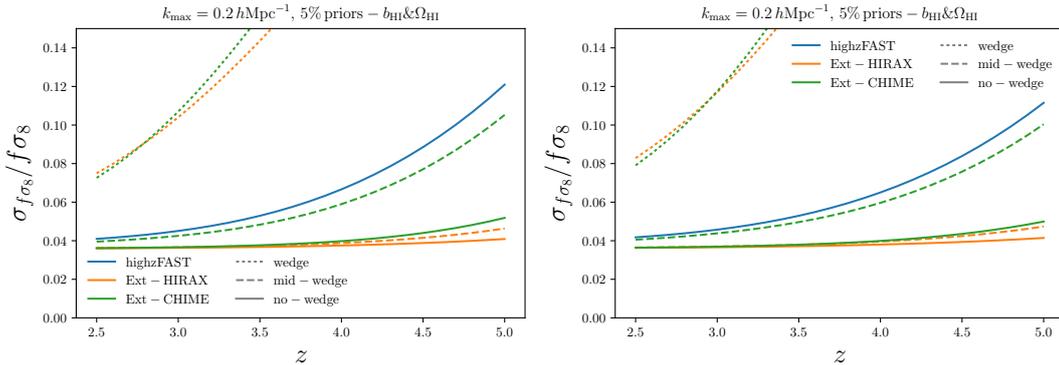


Figure 5. Dependence of the constraints on $f\sigma_8$ on the fiducial value of the HI bias. The left and right panels show the results when the value of M_{\min} , a parameter that controls the cut-off mass in the $M_{\text{HI}}(M, z)$ function and therefore the amplitude of the HI bias, is set to $2 \times 10^{10} M_{\odot}/h$ and $2 \times 10^{11} M_{\odot}/h$, respectively. In both panels we consider a 5% prior on the value of both Ω_{HI} and b_{HI} . Different colors and line types represent the different instruments and assumptions about the wedge (see legend), correspondingly.

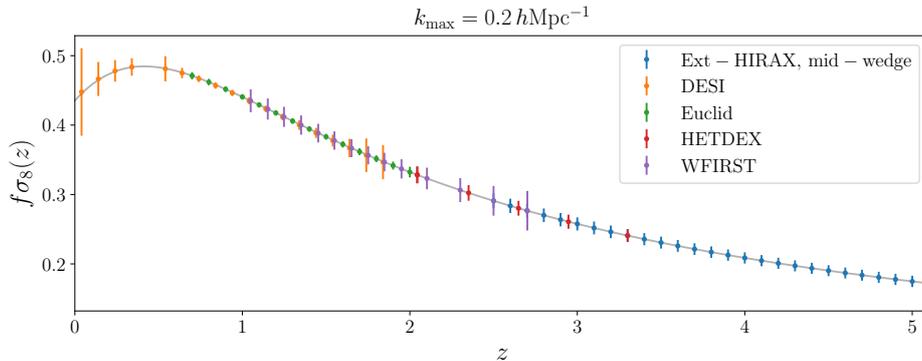


Figure 6. The evolution of $f\sigma_8$ and the predicted constraints as a function of redshift using Ext-HIRAX considering mid wedge configuration with 5% priors on both Ω_{HI} and b_{HI} (blue). For comparison we also show predicted measurements (see text for the references) from upcoming galaxy surveys: DESI (orange, data points shifted by -0.01 in redshift), Euclid (green), HETDEX (red) and WFIRST (purple).

and 2.5 from [23]), Euclid (Table 1.4 from [74], reference case), HETDEX (Figure 6. from [75]), WFIRST (Figure C.6 from [26]) and Ext-HIRAX for the mid-wedge case (Figure 4). We emphasise that the one should not directly compare the precision a survey could possibly achieve in the Figure 6, since different analysis have different underlying assumptions, but the redshift coverage. 21 cm surveys could be able to fill the gap almost to the end of reionization, in a redshift range difficult to reach for conventional surveys. As a reference we used the case of 5% priors and partial cleaning of the wedge.

3.2 BAO distance scale parameters

We now turn our attention to the BAO distance scale parameters. This case is similar to a real analysis if the modelling of the broadband power spectrum has introduced enough

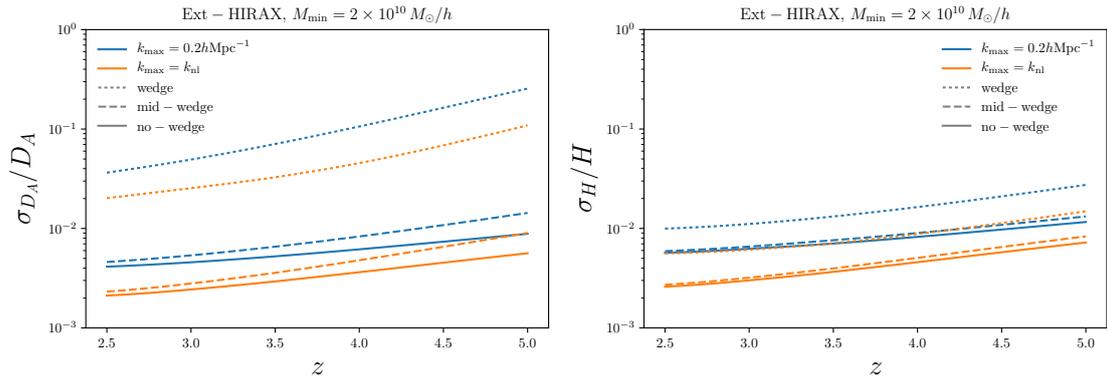


Figure 7. 1σ constraints on the angular diameter distance (left) and Hubble function (right) from the Ext-HIRAX setup. The forecasts have been derived for two different values of $k_{\max} = 0.2 h\text{Mpc}^{-1}$ (blue) and k_{nl} (orange). The solid, dashed and dotted lines represent different assumptions over the wedge (see legend). The fiducial value of the HI bias is computed assuming $M_{\min} = 2 \times 10^{10} M_{\odot}/h$.

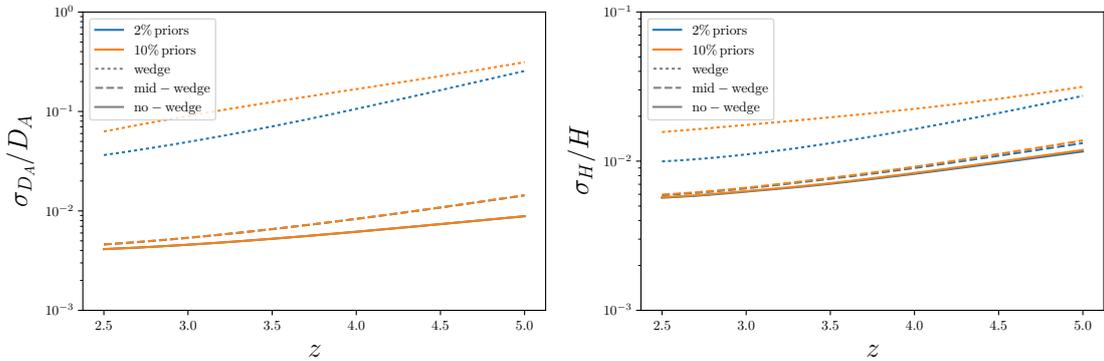


Figure 8. 1σ constraints on the angular diameter distance (left) and Hubble function (right) from the Ext-HIRAX setup assuming different prior knowledge on both b_{HI} and Ω_{HI} : 2% (blue) and 10% (orange). The forecasts have been derived for $k_{\max} = 0.2 h\text{Mpc}^{-1}$. The solid, dashed and dotted lines represent different assumptions over the wedge (see legend). The fiducial value of the HI bias is computed assuming $M_{\min} = 2 \times 10^{10} M_{\odot}/h$.

nuisance parameters that only an amplitude can be measured with enough precision. In this case any constraint on $f\sigma_8$ makes sense only in combination with CMB data.

Results are shown in Figure 7. We find that 21cm surveys can provide very accurate measurements of both the Hubble parameter and the angular diameter distance. For most cases, the geometry of the Universe can be constrained at the sub % precision. Similar to the analysis of the growth of structure in the previous section, the presence of the wedge suppresses the amount of information that can be extracted. The Hubble parameter however can still be measured quite accurately. This is expected since purely radial modes are the only ones surviving the wedge.

The precision at which we can measure the BAO distance scale parameters depends weakly on the prior on astrophysical parameters in the cases of low wedge coverage, while the effect is more pronounced in the full wedge case. We show the 1σ constraints on $H(z)$

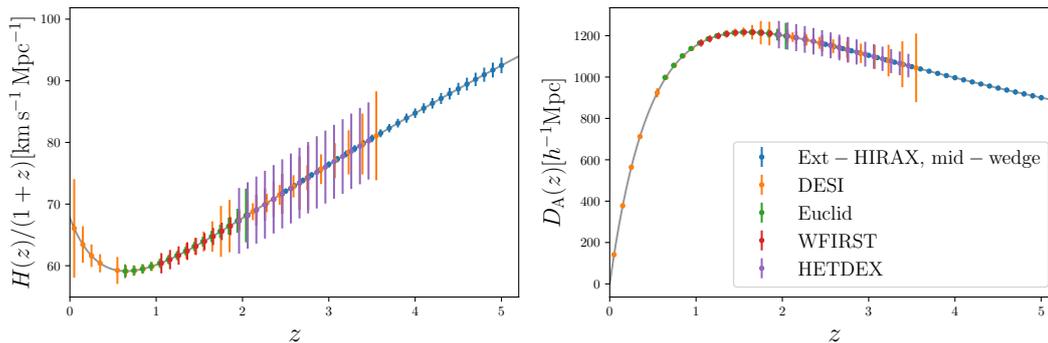


Figure 9. The evolution of the Hubble function $H(z)$ (left panel) and the angular diameter distance $D_A(z)$ (right panel) with the predicted constraints as a function of redshift using Ext-HIRAX considering mid wedge configuration with 5% priors on both Ω_{HI} and b_{HI} (blue). For comparison we also show predicted measurements (see text for the references) from upcoming galaxy surveys: DESI (orange), Euclid (green), HETDEX (red) and WFIRST (purple). Quoted data points have been shifted in redshift by ± 0.01 in order to avoid overlapping.

and $D_A(z)$ dependence on the prior on astrophysical parameters in Figure 8 in the case of Ext-HIRAX.

Similar to figure 6, in Figure 9 we plot the Hubble function $H(z)$ and the angular diameter distance $D_A(z)$ projected measurements from low to high-redshift combining different probes. The forecasted constraints for DESI are taken from [23], while the numbers for other galaxy surveys – Euclid, HETDEX and WFIRST, have all been taken from the corresponding tables in [73]. We also show the forecasted constraints from Ext-HIRAX for the mid-wedge case and assuming 5% priors on b_{HI} and Ω_{HI} . Once again it is clear the contribution of a 21cm survey in filling up the whole redshift range. In Table 3 we list the constraints for both $H(z)$ and D_A as a function of redshift for different foreground configurations of Ext-HIRAX, fixing the 5% priors on b_{HI} and Ω_{HI} .

4 Extension to Λ CDM: Results from probe combination

Cosmological constraints from a single survey usually are limited by the presence of degeneracies between parameters, which can be partially broken by combining different datasets. The purpose of this section is to study the possible benefits of combining data from CMB, galaxy surveys data at low-redshift and 21cm data at high-redshift. Since the three probes we consider do not overlap in the redshift range, the Fisher forecasts become very simple as the covariance is diagonal and the different Fisher matrices can be added up.

We include the current constraints on the cosmological parameters measured from Planck 2015 [1] in combination with BAO measurements from BOSS at redshift $z \simeq 0.3$ and $z \simeq 0.5$ [76]. In order to obtain the corresponding Fisher matrix we use the publicly available MCMC chains⁷ which are based on Planck temperature (TT) + low TEB polarisation + external BAO constraints. We then compute the covariance matrix for the parameters we consider and invert it.

⁷We use `base_mnu_plikHM_TT_lowTEB_BAO` chains for the study on the sum of neutrino masses Σm_ν and `base_nnu_plikHM_TT_lowTEB_BAO` chains for the study of the number of relativistic species N_{eff} .

Additionally, we use the Fisher matrix calculated for a generic CMB Stage 4 experiment⁸. We refer the reader to the CMB S4 Science Book [77] for more details about next generation of CMB experiments. For consistency we remove Planck information if CMB S-4 information is used.

We also separately perform Fisher forecasts for future spectroscopic galaxy redshift surveys such as Euclid [24] or DESI [73], rather than relying on existing results. This choice has been made in order to avoid the final constraints to be affected by different modeling assumptions in different analysis. It is also motivated by the need for the full Fisher matrix rather than the marginalized error on each individual parameter.

To model the galaxy power spectrum, we employ a similar expression to eq. (2.26), but without the brightness temperature factor and using the galaxy bias instead of the HI bias. The effect of non-linearities on the galaxy power spectrum is larger at lower redshifts and effectively damps the BAO oscillations in the power spectrum. We use the following description to model this effect in redshift-space [37]:

$$P_{\text{nl}}^{\text{RSD}}(k, \mu) \propto P_{\text{lin}}^{\text{RSD}}(k, \mu) \exp\left(-\frac{k_{\perp}^2 \Sigma_{\perp}^2}{2} - \frac{k_{\parallel}^2 \Sigma_{\parallel}^2}{2}\right), \quad (4.1)$$

where the term $P_{\text{lin}}^{\text{RSD}}(k, \mu)$ contains the bias and linear redshift-space distortions terms, while Σ_{\perp} and Σ_{\parallel} are the non-linear damping scales in the transverse and the radial direction, respectively. We compute the damping scales using $\Sigma_{\perp} = 9.4(\sigma_8(z)/0.9) h^{-1} \text{Mpc}$ and $\Sigma_{\parallel} = (1 + f(z))\Sigma_{\perp}$ [73]. This effect is particularly important for the BAO measurements at low-redshifts since it deteriorates the constraints on $D_A(z)$ and $H(z)$. In our forecasts for the massive neutrinos we will not assume the BAO reconstruction (see Section 2.1) has been applied and we will use the full values of Σ_{\perp} and Σ_{\parallel} , while in our forecasts on the N_{eff} we do assume BAO reconstruction and half the damping scales accordingly. The exponential factor in eq. (4.1) is taken outside the logarithmic derivatives in the Fisher matrix since we do not want to infer any distance information from these factors. In practice, the exponential only degrades the effective volume.

For Euclid, we assume an area equal to 15000 deg^2 covering the redshift range $0.65 < z < 2.05$. We take the fiducial galaxy bias $b(z) = \sqrt{1+z}$ and treat it as a free parameter that we marginalise over in each redshift bin, $\Delta z = 0.1$. The only noise term in the power spectrum we consider is the shot-noise coming from the discreteness of galaxy sample and equal to the inverse number density of galaxies: $P_{\text{N}}^{\text{shot}} = \bar{n}^{-1} [h^{-3} \text{Mpc}^3]$. We compute the \bar{n} using the data in Table 1.3 (reference case) of [74]. We will show the results for two different values of the maximum wavenumber: Euclid02 for which we take $k_{\text{max}} = 0.2 h \text{Mpc}^{-1}$ and Euclidnl for which we take $k_{\text{max}} = k_{\text{nl}}(z)$.

Of all possible extension of the standard cosmological model we focus on the sum of the neutrino masses, $\sum m_{\nu}$, and on the effective number of relativistic degrees of freedom, N_{eff} , since, as we will see, these two have a clear target to be reached.

4.1 Massive neutrinos

Constraining neutrino masses is one of the main goals of current and future cosmological surveys. Massive neutrinos suppress power on intermediate and small scales and can therefore be detected by combining probes of large scales, like CMB anisotropies, with data from intermediate/small scales [78]. Currently the best constraints are $\sum m_{\nu} < 0.11 \text{ eV}$ [79], but

⁸We thank Anže Slosar for providing us the Fisher matrix of CMB S-4 made by Joel Meyers.

we expect the next generation of galaxy surveys and CMB probes to significantly improve those bounds. The benchmark to achieve for the error on neutrino masses is set by the laboratory measurements of neutrino oscillations. These indicate two possible scenarios for the sum of neutrino masses, the normal hierarchy where $\sum m_\nu \geq 0.06$ eV and the inverted one where $\sum m_\nu \geq 0.1$ eV [80].

The goal of this section is to study the improvements on the constraints on the neutrino masses that can be achieved by adding information from 21cm intensity mapping surveys in the redshift range $2.5 < z < 5$. Improvements on the constraints coming from the epoch of reionization have been presented in [81–83].

In the case of massive neutrinos we have to slightly modify our signal model. In fact, as shown [84–87], halos and galaxies are biased tracers of the CDM+baryons field only, as opposed to the total matter field – CDM+baryons+neutrinos. This means that we have to change the $P(k)$ and $f(k)$ in eq. (2.1) to include only the CDM+baryons component. This important physical effect has the consequence of reducing the constraining power of a broadband analysis of $P_{21}(k, z)$, since the total matter power spectrum is more suppressed by the presence of massive neutrinos than the CDM+baryons only. In our case, we find the difference being around 30% in terms of the forecasted σ_{m_ν} , an absolutely non-negligible effect. To the best of our knowledge, existing forecasts on neutrino masses do not include this effect in their analysis (with some notable exceptions as in [12]), suggesting some level of bias in their final answers. Our forecast is therefore more realistic, but it yields worse constraints.

Another important thing to notice is that the effect of neutrinos on the power spectrum become smaller with redshift. This is due to the fact that neutrinos have less time to delay the growth of CDM+baryons perturbations on small scales. This means that the constraint on Σm_ν from two surveys, one at low-redshift and another at high-redshift, covering the same volume and using the same amount of information ($k_{\min} < k < k_{\max}$) will be different, the high-redshift one being worse. In contrast, going to higher redshifts has two advantages: larger survey volume and larger k_{nl} up to which one can hope to use perturbation theory. Figure 10 shows the ratio between the linear CDM+baryons (dashed) and the total matter (solid) power spectrum in a model with massive neutrinos ($\Sigma m_\nu = 0.06$ eV) to the same quantities in the standard Λ CDM model ($\Sigma m_\nu = 0$ eV). Different colors represent different redshift, in blue $z = 1$ and in red $z = 4$. We notice that the difference between CDM+baryons and total matter remains constant with redshift, but the overall suppression compared to the vanilla models decrease by almost a factor of two between low and high z .

We show the forecasted constraints on the neutrino masses and other cosmological parameters coming from external datasets alone in table 4. Constraints on the neutrino masses for various combinations of probes and different foreground assumptions are shown in Table 7 and 8 for Ext-CHIME and Ext-HIRAX, respectively. We see that in the most pessimistic case, i.e. full wedge and $k_{\max} = 0.2 h \text{ Mpc}^{-1}$, Ext-HIRAX plus Euclid and CMB-S4 is able to pin down neutrino masses to $\sigma_{m_\nu} = 0.028$ eV, with intensity mapping bringing a 10% improvement on the errorbars. In the most optimistic case this combination of probes yields $\sigma_{m_\nu} = 0.018$ eV in which case the gain of adding IM can reach 40%. It is important to stress again that the use of CDM+baryons to describe the clustering of galaxies and 21cm is worsening the error bars by roughly 30%. In Figure 11 we show the constraints and the degeneracies of the neutrino masses with other cosmological parameters when Ext-HIRAX is combined with CMB-S4 and CMB-S4+Euclid in the most optimistic case. Single-dish experiments do not help in constraining neutrinos masses as well as interferometers, as one can see from results presented in Tables 6 for highzFAST alone and Table 9 in which we show the

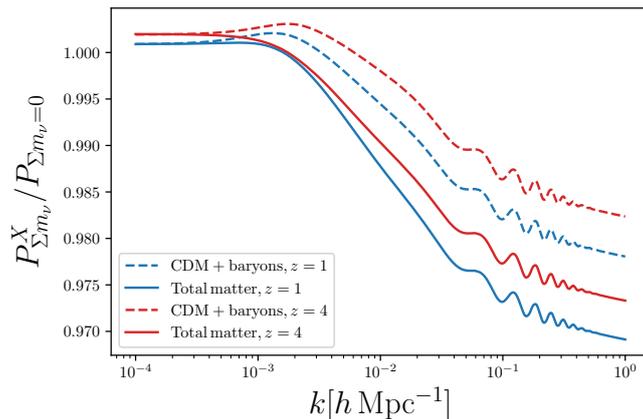


Figure 10. Ratio between the linear CDM+baryons (dashed) and the total matter (solid) power spectrum in a model with massive neutrinos ($\Sigma m_\nu = 0.06$ eV) to the same quantities in the standard Λ CDM model ($\Sigma m_\nu = 0$ eV). Different colors represent different redshift, in blue $z = 1$ and in red $z = 4$.

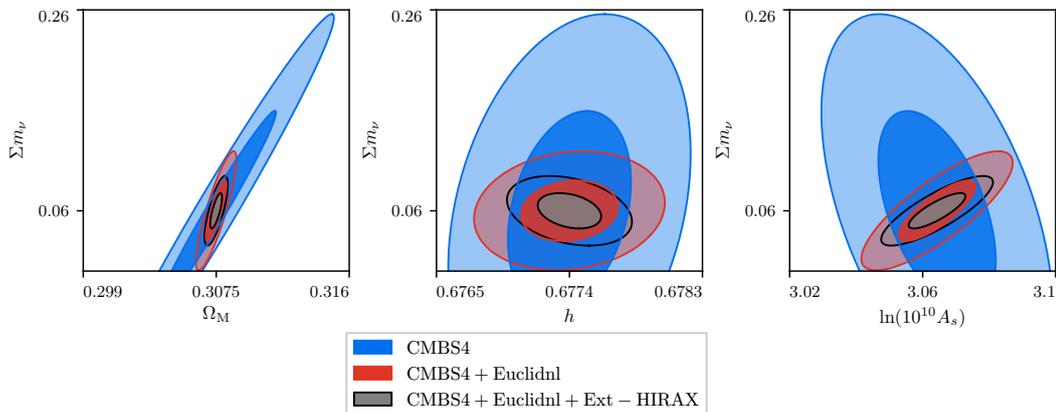


Figure 11. Constraints on the neutrino masses and their degeneracies with Ω_m (left), h (middle) and A_s (right) from CMBS4 (blue), CMBS4+Euclid (red) and CMBS4+Euclid+Ext-HIRAX (black). For Ext-HIRAX we have assumed $k_{\max} = k_{\text{nl}}$, no wedge and 2% priors on both b_{HI} and Ω_{HI} .

results when combining different probes together.

We notice that the effect of the wedge does not worsen the constraints too much. In Section 3.1 we have shown that the presence of the wedge deteriorate significantly the constraints on $f\sigma_8$. Thus, we conclude that most of the improvement brought in by the 21cm data in the combined analysis is due to the very accurate distance measurements, which are less affected by foregrounds. A similar argument applies to external priors on the density and the bias of the HI field, which do not change much the final result on neutrino masses. On the other hand, an even more aggressive, in terms of signal-to-noise, instrument could potentially improve on neutrino masses constraints.

Another advantage of observing the high-redshift Universe is that dark energy is a sub dominant fraction of the energy density. It is in fact well known that in the extended parameter spaces the equation of state of dark energy w is very degenerate with other parameters, e.g. massive neutrinos and curvature [88]. In Figure 12 we show the degeneracy $\Sigma m_\nu - w$

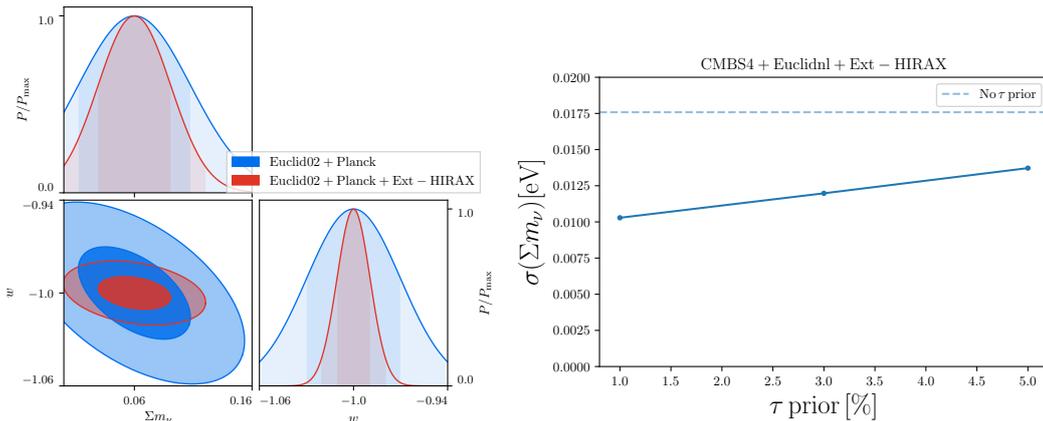


Figure 12. (Left panel) Degeneracy between Σm_ν and w when considering only Euclid02+Planck (blue) and when adding information from Ext-HIRAX (red). (Right panel) The effect of an additional τ prior on the neutrino masses constraints coming from the combination of CMBS4 + Euclidnl + Ext-HIRAX. The blue dashed line shows the case on no additional prior assumed.

degeneracy in the case where only low redshift measurements in combination with CMB are available (Euclid02+Planck) and when the information from above $z > 2.5$ (Ext-HIRAX) is added on top. We find a factor of 3 improvement on the error on w when adding Ext-HIRAX, constraining w to the 1% level even if neutrinos are allowed to vary. For Ext-HIRAX in this case we have used $k_{\max} = 0.2 \text{ h Mpc}^{-1}$, 2% priors on both b_{HI} and Ω_{HI} , and mid-wedge case.

To conclude this Section we would like to emphasize that even adding high-redshift information the constraint on neutrino masses is primarily limited by the uncertainty in the amplitude of the primordial fluctuation spectra A_s and the optical depth of reionization – τ – in CMB data. The parameter τ is weakly constrained by the current Planck results [1], with an error bar almost a factor of two worse than the forecasted value [89]. Improvement in the understanding of systematics in the Planck polarization data will likely improve the constraint on τ , but to be more conservative we decided to adopt the current estimate. This is not the case for pre-Planck forecasts, e.g. DESI [73] or more recent ones [23, 90], who assume in their forecasts the blue book value for σ_τ . This fact, together with our correct choice of the modeling at the beginning of this Section, explains why our constraint on neutrino masses is in general worse than others in the literature. Concerning the future improvements on constraining τ , aside from including all polarization measurements from Planck, another hope is that future reionization epoch IM surveys, e.g. HERA, could bring tight constraints on τ [91], thereby bringing substantial improvements on other cosmological parameters. Figure 12 shows how much a better measurement of τ in the CMB could help in constraining the neutrino masses in the best case of Ext-HIRAX with $k_{\max} = k_{\text{nl}}$, 2% prior on both b_{HI} and Ω_{HI} , and no wedge.

4.1.1 Beyond linear theory

So far we have considered a simple model for the 21cm power spectrum consisting of linear theory and linear bias term. Going beyond linear theory requires new nuisance parameters that need to be marginalised over [71, 72, 90, 92]. This can degrade the constraints on the neutrino masses. In order to estimate the effect of these extra terms we will consider a simple order-of-magnitude extension to our power spectrum model and write the 21cm power

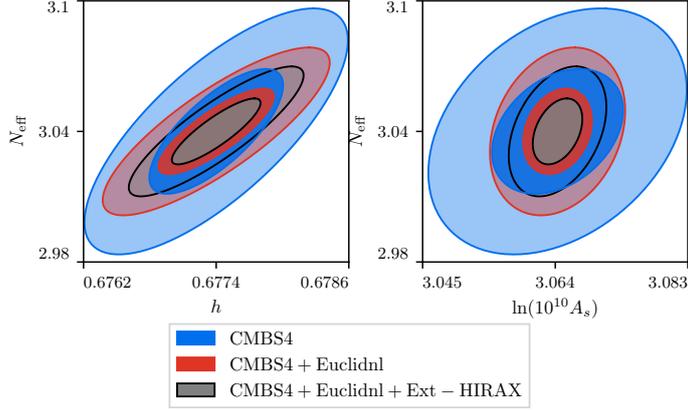


Figure 13. Constraints on N_{eff} and its degeneracies with h (left) and A_s (right) from CMB-S4 (blue), CMB-S4+Euclid (red) and CMB-S4+Euclid+Ext-HIRAX (black). For Ext-HIRAX we have assumed $k_{\text{max}} = k_{\text{nl}}$, no wedge and 2% priors on both b_{HI} and Ω_{HI} . We used `cosmicfish` software [95] to make this Figure and also Figure 11.

spectrum as:

$$P_{\text{HI}}(k, z) = \bar{T}_b^2 (b_{\text{HI}}^2 + f\mu^2)^2 P_{\text{lin}}(k, z=0) D(z)^2 \left(1 + 2k^2 R^2 + \alpha P_{\text{lin}}(k, z=0) D(z)^2 \frac{k^3}{2\pi^2} \right). \quad (4.2)$$

The second term in the parentheses has the form of the most relevant counter-term in the effective field theory approach to perturbation theory [93, 94]. The third term is included in order to mimic the effects of 1-loop corrections to the power spectrum and can be understood as the “theoretical error” [72] on the linear theory model. This approximation for one-loop contributions is a good order-of-magnitude estimate on scales relevant for constraining the sum of the neutrino masses. The coefficients R and α are free parameters in each redshift bin and we take the fiducial values $R = \alpha = 0$. We marginalise over R and α in each redshift bin. We show the resulting constraints on the cosmological parameters from Ext-HIRAX in Table 5. Compared to the standard case of linear theory and bias, the constraints on the neutrino masses become weaker by a factor of $\sim 25\%$ and 31% in the no wedge case of $k_{\text{max}} = 0.2 h \text{ Mpc}^{-1}$ and $k_{\text{max}} = k_{\text{nl}}(z)$, respectively. Adding other probes (Euclid, CMB S-4) makes this difference smaller and the constraints are weaker by a factor of 5–10% (see Table 7). This reinforces our previous argument that most of the constraining power is coming from geometrical measurements of the BAO distances.

4.2 The effective numbers of relativistic degrees of freedom

Another extension of the standard ΛCDM model is represented by the presence of extra radiation species. In the standard model, with 3 massive neutrinos, one has $N_{\text{eff}} = 3.046$ [96], but any light species in thermal equilibrium during the history of the Universe will increase this value. Theoretical estimates in [97, 98] set the minimum ΔN_{eff} due to any light particles between $z = 0$ and the QCD phase transition to $\Delta N_{\text{eff}} = 0.027$. This number is therefore the natural target of any analysis of N_{eff} .

In CMB data, N_{eff} is very degenerate with other parameters, and the 1σ forecasted error of CMB S-4 is slightly above the theory benchmark. However, it has been recently realized that a clean signature of N_{eff} exists and can be detected in both CMB and large

scale structure [99–101], as the presence of extra relativistic degrees of freedom introduce a phase shift in the acoustic peaks of the baryon-photon fluid [99]. This effect has already been detected in Planck data [100, 102]. In [101], the authors have shown that this phase shift survives non-linear evolution and therefore could in principle be detected in BAO analysis of galaxy clustering.

Building and testing a template for the phase shift goes beyond the scope of this paper, but motivated by the above discussion we study the constraints on N_{eff} from various combination of probes. We fit the full broadband power spectrum, which by definition contains also the phase shift, to constraint the number of relativistic degrees of freedom. In a realistic data analysis one would use reconstructed data, which we mimic by dividing by a factor of 2 the damping factor in eq. (4.1) [73].

We show the forecasted constraints on N_{eff} and other cosmological parameters coming from external datasets alone in table 10. Our results for various combinations of probes and different foreground assumptions are presented in tables 11, 12, 13, 14, 15 and figure 13. Our forecasted error for Ext-HIRAX ranges from $\sigma_{N_{\text{eff}}} = 0.015$, for the very optimistic case of no wedge and $k_{\text{max}} = k_{\text{nl}}$, to $\sigma_{N_{\text{eff}}} = 0.021$ for the full wedge and $k_{\text{max}} = 0.2 h \text{ Mpc}^{-1}$ case. These numbers are interesting, but on the other hand any deviation from the minimal value will be detected at less than 1σ . It is also important to point out that 21cm surveys will add very little to the combination of CMB data and galaxy surveys. Our results are in broad agreement with the analysis of [101].

5 Summary and conclusions

In this work we have explored the possibility of studying cosmology through radio-telescopes that operate in the redshift range $2.5 < z < 5$. The reason behind this is that, while this is a redshift-range not considered in current and upcoming setups, the volume it encloses is much larger than the one probed by current and future spectroscopic surveys. The question we try to answer is: how much cosmological information is contained in this redshift window?

We focus our analysis on four key cosmological quantities: 1) the growth rate, $f\sigma_8$, the BAO distance scale parameters, D_A and H , the sum of the neutrino masses, Σm_ν , and the number of relativistic degrees of freedom, N_{eff} . We consider four extensions of current or upcoming radio-telescopes like HIRAX, CHIME and FAST, and two different observational strategies: interferometry (for Ext-HIRAX and Ext-CHIME) and single-dish (highzFast).

We carry out our analysis using the Fisher matrix formalism. We model the amplitude and shape of the 21cm signal using the model proposed in [48]. We also account for cosmological and instrumental effects such as the presence of the wedge, the window function, the instrument thermal noise, the angular resolution, the presence of shot-noise...etc.

We point out that measurements that are sensitive to the overall amplitude of the 21cm power spectrum, like $f\sigma_8$, will be completely degenerate with astrophysical parameters like Ω_{HI} and b_{HI} . In order to break that degeneracy it is necessary that we use independent datasets that constrain those quantities. We show how the value of these parameters can be determined through either the Ly α -forest alone or via cross-correlations between 21cm and the Ly α -forest or DLAs.

Under the assumption of the primary beam foreground wedge contamination (mid-wedge case in the text), 5% priors on b_{HI} and Ω_{HI} and $k_{\text{max}} = 0.2 h \text{ Mpc}^{-1}$, that we term the fiducial setup, we find that Ext-HIRAX can constrain the value of $f\sigma_8$ within bins of $\Delta z = 0.1$ at $\simeq 4\%$ in the redshift range $2.5 < z < 5$. A modest improvement is achieved by changing k_{max}

from $0.2 h\text{Mpc}^{-1}$ to k_{nl} . If data from the whole wedge need to be discarded, these constraints degrade between a factor 2 (at $z = 2.5$) and 7 (at $z = 5$).

Under the fiducial setup, we find that Ext-HIRAX will place $\simeq 1\%$ constraints on D_A and H . As with the growth rate, our results point out that going to smaller scales has only a very modest impact on the results. Being able to use a fraction of the modes in the wedge has a huge impact on our results, as removing the information in the whole wedge degrades the constraints between a factor of 10 (at $z = 2.5$) and 20 (at $z = 5$).

We have also studied the impact that the theory model has on the results. By using a theory template that accounts for 1-loop corrections and incorporates 2 free parameters that we marginalise over, we find that the constraints on the cosmological parameters worsen between 10% and 500% when 2% priors on Ω_{HI} and b_{HI} are used. In the case of the neutrino masses, the constraints worsen between 10% and 30%.

We find that data from Ext-HIRAX in the fiducial setup can constrain the neutrino masses with an error of 0.10 eV. In combination with data from CMB S4 and galaxy clustering from Euclid the errors shrink to $\simeq 20$ meV. Our results are not very sensitive to the wedge coverage, the minimum scale employed and the priors on b_{HI} and Ω_{HI} .

Finally, we find that data from Ext-HIRAX, in the fiducial setup, plus CMB S4 plus Euclid can constrain N_{eff} with a very competitive error of 0.02. As with the neutrino masses, our constraints do not depend much on the Ω_{HI} and b_{HI} priors, k_{max} and the wedge coverage.

Results for the Ext-CHIME and highzFAST instruments are similar to those of Ext-HIRAX, with the exception of neutrino masses, where highzFAST performs worse than Ext-CHIME or Ext-HIRAX.

We conclude that there is a large amount of cosmological information embedded in the, poorly constrained, redshift range $2.5 < z < 5$. Suitable extensions of existing and upcoming radio-telescopes targeting at this redshift window can provide very tight constraint on key cosmological parameters.

6 Tables

In this section we show the tables with the constraints on RSD, BAO distance scale parameters and cosmological parameters coming from various configurations of the proposed 21cm IM surveys or in combination with current and foreseen galaxy surveys and CMB experiments.

Ext-HIRAX $\sigma_{f\sigma_8}/(f\sigma_8)$, 5% b_{HI} & Ω_{HI}						
z	$k_{\text{max}} = 0.2h\text{Mpc}^{-1}$			$k_{\text{max}} = k_{\text{nl}}(z)$		
	No wedge	Mid wedge	Wedge	No wedge	Mid wedge	Wedge
2.5	0.036	0.036	0.075	0.035	0.035	0.043
2.6	0.036	0.036	0.080	0.035	0.036	0.044
2.7	0.036	0.036	0.085	0.035	0.036	0.046
2.8	0.036	0.036	0.091	0.035	0.036	0.048
2.9	0.036	0.037	0.097	0.036	0.036	0.050
3.0	0.036	0.037	0.104	0.036	0.036	0.053
3.1	0.036	0.037	0.111	0.036	0.036	0.056
3.2	0.036	0.037	0.119	0.036	0.036	0.060
3.3	0.037	0.037	0.127	0.036	0.036	0.065
3.4	0.037	0.037	0.135	0.036	0.036	0.070
3.5	0.037	0.037	0.144	0.036	0.036	0.075
3.6	0.037	0.038	0.153	0.036	0.036	0.082
3.7	0.037	0.038	0.162	0.036	0.036	0.089
3.8	0.037	0.038	0.171	0.036	0.036	0.097
3.9	0.037	0.038	0.181	0.036	0.036	0.105
4.0	0.038	0.039	0.190	0.036	0.036	0.114
4.1	0.038	0.039	0.200	0.036	0.036	0.124
4.2	0.038	0.040	0.209	0.036	0.037	0.135
4.3	0.038	0.040	0.218	0.036	0.037	0.146
4.4	0.039	0.041	0.228	0.036	0.037	0.158
4.5	0.039	0.041	0.237	0.036	0.037	0.170
4.6	0.039	0.042	0.246	0.037	0.038	0.182
4.7	0.040	0.043	0.254	0.037	0.038	0.194
4.8	0.040	0.044	0.263	0.037	0.038	0.206
4.9	0.040	0.045	0.271	0.037	0.039	0.219
5.0	0.041	0.046	0.279	0.037	0.040	0.231

Table 2. Forecasted 1σ constraints on $f\sigma_8$ for Ext-HIRAX as a function of redshift for different wedge configurations and different k_{max} . The numbers correspond to 5% priors on the values of both Ω_{HI} and b_{HI} .

Ext-HIRAX, $k_{\max} = 0.2h\text{Mpc}^{-1}$, 5% b_{HI} & Ω_{HI} ,						
z	σ_{D_A}/D_A			σ_H/H		
	No wedge	Mid wedge	Wedge	No wedge	Mid wedge	Wedge
2.5	0.004	0.005	0.053	0.006	0.006	0.014
2.6	0.004	0.005	0.057	0.006	0.006	0.014
2.7	0.004	0.005	0.060	0.006	0.006	0.014
2.8	0.004	0.005	0.064	0.006	0.006	0.014
2.9	0.004	0.005	0.067	0.006	0.006	0.014
3.0	0.005	0.005	0.071	0.006	0.007	0.014
3.1	0.005	0.006	0.075	0.006	0.007	0.015
3.2	0.005	0.006	0.080	0.007	0.007	0.015
3.3	0.005	0.006	0.084	0.007	0.007	0.015
3.4	0.005	0.006	0.089	0.007	0.007	0.016
3.5	0.005	0.007	0.095	0.007	0.008	0.016
3.6	0.005	0.007	0.10	0.007	0.008	0.016
3.7	0.006	0.007	0.11	0.008	0.008	0.017
3.8	0.006	0.008	0.11	0.008	0.009	0.017
3.9	0.006	0.008	0.12	0.008	0.009	0.018
4.0	0.006	0.008	0.13	0.008	0.009	0.019
4.1	0.006	0.009	0.14	0.009	0.010	0.019
4.2	0.007	0.009	0.15	0.009	0.010	0.020
4.3	0.007	0.010	0.16	0.009	0.010	0.021
4.4	0.007	0.010	0.17	0.010	0.011	0.022
4.5	0.007	0.011	0.18	0.010	0.011	0.023
4.6	0.008	0.011	0.20	0.010	0.012	0.024
4.7	0.008	0.012	0.21	0.011	0.012	0.025
4.8	0.008	0.013	0.23	0.011	0.013	0.026
4.9	0.009	0.014	0.25	0.011	0.013	0.027
5.0	0.009	0.014	0.27	0.012	0.014	0.029

Table 3. Forecasted 1σ constraints on the expansion rate $H(z)$ and the angular diameter distance $D_A(z)$ for Ext-HIRAX as a function of redshift for different wedge configurations. The numbers correspond to 5% priors on the values of both Ω_{HI} and b_{HI} and $k_{\max} = 0.2h\text{Mpc}^{-1}$.

	External datasets – Σm_ν				
	Ω_{M}	h	$\Sigma m_\nu[\text{eV}]$	$\ln(10^{10} A_s)$	n_s
Fiducial values	0.3075	0.6774	0.060	3.064	0.9667
PlanckBAO	0.0082	0.0065	0.067	0.037	0.0048
Euclid02 No Rec.	0.0055	0.0025	0.20	0.091	0.0086
Euclidnl No Rec.	0.0050	0.0021	0.17	0.077	0.0065
CMB-S4	0.0038	0.0004	0.10	0.018	0.0019

Table 4. Forecasted 1σ constraints on cosmological parameters and the Σm_ν considering each external dataset alone. These constraints are obtained using the Fisher matrices explained in Section 4. PlanckBAO represents Planck CMB combined with BOSS BAO measurements. We show the projected constraints for Euclid for two different k_{\max} used and the constraints coming from CMB-S4 experiment.

Ext-HIRAX Σm_ν					
	Ω_M	h	Σm_ν [eV]	$\ln(10^{10} A_s)$	n_s
Fiducial values	0.3075	0.6774	0.060	3.064	0.9667
No wedge $k_{\max} = 0.2h\text{Mpc}^{-1}$					
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0016	0.0010	0.059	0.026	0.0036
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$, diff M_{\min}	0.0015	0.0009	0.056	0.025	0.0033
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$, 1-loop	0.0020	0.0010	0.081	0.038	0.014
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0024	0.0015	0.093	0.047	0.0038
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0029	0.0018	0.11	0.065	0.0040
No wedge $k_{\max} = k_{\text{nl}}(z)$					
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0010	0.0007	0.040	0.020	0.0016
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$, diff M_{\min}	0.0009	0.0006	0.037	0.019	0.0014
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$, 1-loop	0.0011	0.0007	0.058	0.027	0.0060
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0014	0.0009	0.054	0.032	0.0016
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0014	0.0009	0.054	0.032	0.0016
Mid wedge $k_{\max} = 0.2h\text{Mpc}^{-1}$					
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0017	0.0010	0.062	0.027	0.0039
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$, diff M_{\min}	0.0016	0.0010	0.060	0.027	0.0037
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$, 1-loop	0.0021	0.0011	0.088	0.041	0.016
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0025	0.0016	0.098	0.049	0.0042
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0032	0.0020	0.12	0.071	0.0044
Mid wedge $k_{\max} = k_{\text{nl}}(z)$					
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0011	0.0007	0.042	0.021	0.0017
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$, diff M_{\min}	0.0010	0.0006	0.039	0.020	0.0015
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$, 1-loop	0.0011	0.0008	0.061	0.028	0.0066
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0015	0.0010	0.058	0.035	0.0017
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0017	0.0011	0.065	0.051	0.0018
Wedge $k_{\max} = 0.2h\text{Mpc}^{-1}$					
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0066	0.0048	0.24	0.091	0.010
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$, diff M_{\min}	0.0068	0.0050	0.25	0.092	0.010
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$, 1-loop	0.0076	0.0050	0.34	0.138	0.054
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0072	0.0053	0.27	0.107	0.011
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0084	0.0062	0.31	0.136	0.011
Wedge $k_{\max} = k_{\text{nl}}(z)$					
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0027	0.0019	0.11	0.042	0.0044
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$, diff M_{\min}	0.0026	0.0019	0.10	0.041	0.0040
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$, 1-loop	0.0029	0.0021	0.16	0.060	0.018
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0034	0.0025	0.13	0.059	0.0044
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0048	0.0034	0.17	0.092	0.0045

Table 5. Fiducial values and 68% confidence intervals on cosmological parameters and the Σm_ν using Ext-HIRAX alone for various considerations. Going from top to bottom, we show the constraints considering different priors on both $b_{\text{HI}} \& \Omega_{\text{HI}}$, different k_{\max} we use and different level of foreground wedge contamination. Alongside, focusing only on the case of 2% priors on $b_{\text{HI}} \& \Omega_{\text{HI}}$, we show the results obtained when we use a different value of $M_{\min} = 2 \times 10^{11} M_\odot/h$ (labeled diff M_{\min} .) and when we go beyond linear theory and marginalise over 1-loop and counter-terms parameters (labeled 1-loop, see Section 4.1.1).

highzFAST Σm_ν					
	Ω_M	h	Σm_ν [eV]	$\ln(10^{10} A_s)$	n_s
Fiducial values	0.3075	0.6774	0.060	3.064	0.9667
$k_{\max} = 0.2h\text{Mpc}^{-1}$					
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0045	0.0026	0.14	0.054	0.013
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0054	0.0032	0.18	0.076	0.013
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0069	0.0041	0.24	0.112	0.013
$k_{\max} = k_{\text{nl}}(z)$					
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0031	0.0021	0.11	0.044	0.0076
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0041	0.0026	0.15	0.064	0.0078
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0054	0.0034	0.19	0.095	0.0079

Table 6. Fiducial values and 68% confidence intervals on cosmological parameters and the Σm_ν using highzFAST alone for various considerations. We show the constraints considering different astrophysical priors on both $b_{\text{HI}} \& \Omega_{\text{HI}}$, as well as the dependence of the constraints on the k_{\max} used.

Ext-HIRAX Σm_ν [eV]						
	$k_{\max} = 0.2h\text{Mpc}^{-1}$			$k_{\max} = k_{\text{nl}}(z)$		
Euclid No Rec.+PlanckBAO	0.050			0.049		
+21cm	No wedge	Mid wedge	Wedge	No wedge	Mid wedge	Wedge
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.037	0.038	0.045	0.030	0.031	0.040
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$, diff M_{\min}	0.037	0.038	0.044	0.028	0.029	0.039
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$, 1-loop	0.038	0.039	0.046	0.035	0.035	0.042
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.042	0.043	0.048	0.034	0.035	0.043
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.044	0.045	0.049	0.035	0.036	0.044
Euclid No Rec.+CMB-S4	0.031			0.030		
+21cm	No wedge	Mid wedge	Wedge	No wedge	Mid wedge	Wedge
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.022	0.022	0.028	0.018	0.018	0.025
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$, diff M_{\min}	0.021	0.021	0.028	0.017	0.017	0.024
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$, 1-loop	0.023	0.023	0.028	0.020	0.020	0.025
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.023	0.023	0.028	0.018	0.019	0.025
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.023	0.023	0.028	0.019	0.019	0.025

Table 7. Forecasted 1σ constraints on Σm_ν considering Ext-HIRAX instrument combined with CMB+galaxy probes. We show the constraints considering different priors on both $b_{\text{HI}} \& \Omega_{\text{HI}}$, different level of wedge contamination and different k_{\max} used. In the top half of the table we show what are the gains of adding Ext-HIRAX measurements on top of combined Fisher matrix for Euclid galaxy survey and PlanckBAO. In the bottom half we do the same only using CMBS-4 Fisher matrix instead of the PlanckBAO one. Additionally, focusing only on the case of 2% priors on $b_{\text{HI}} \& \Omega_{\text{HI}}$, we also show the results obtained when we use a different value of $M_{\min} = 2 \times 10^{11} M_\odot/h$ (labeled diff M_{\min} .) and when we go beyond linear theory and marginalise over 1-loop and counter-terms parameters (labeled 1-loop, see Section 4.1.1).

Ext-CHIME Σm_ν [eV]						
	$k_{\max} = 0.2h\text{Mpc}^{-1}$			$k_{\max} = k_{\text{nl}}(z)$		
Euclid No Rec.+PlanckBAO	0.050			0.049		
+21cm	No wedge	Mid wedge	Wedge	No wedge	Mid wedge	Wedge
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.038	0.041	0.044	0.033	0.035	0.040
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.043	0.045	0.048	0.037	0.040	0.043
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.044	0.047	0.049	0.037	0.041	0.044
Euclid No Rec.+CMB-S4	0.031			0.030		
+21cm	No wedge	Mid wedge	Wedge	No wedge	Mid wedge	Wedge
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.023	0.026	0.029	0.020	0.022	0.026
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.024	0.026	0.029	0.021	0.023	0.026
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.025	0.026	0.029	0.021	0.023	0.027

Table 8. Forecasted 1σ constraints on Σm_ν considering Ext-CHIME instrument combined with CMB+galaxy probes. We show the constraints considering different priors on both $b_{\text{HI}} \& \Omega_{\text{HI}}$, different level of wedge contamination and different k_{\max} used. In the top half of the table we show what are the gains of adding Ext-CHIME measurements on top of combined Fisher matrix for Euclid galaxy survey and PlanckBAO. In the bottom half we do the same only using CMBS-4 Fisher matrix instead of the PlanckBAO one.

highzFAST Σm_ν [eV]		
	$k_{\max} = 0.2h\text{Mpc}^{-1}$	$k_{\max} = k_{\text{nl}}(z)$
Euclid No Rec.+PlanckBAO	0.050	0.049
+21cm		
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.043	0.041
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.048	0.045
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.049	0.047
Euclid No Rec.+CMB-S4	0.031	0.030
+21cm		
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.029	0.028
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.030	0.028
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.030	0.029

Table 9. Forecasted 1σ constraints on Σm_ν considering highzFAST instrument combined with CMB+galaxy probes. We show the constraints considering different priors on both $b_{\text{HI}} \& \Omega_{\text{HI}}$ and different k_{\max} used. In the top half of the table we show what are the gains of adding highzFAST measurements on top of combined Fisher matrix for Euclid galaxy survey and PlanckBAO. In the bottom half we do the same only using CMBS-4 Fisher matrix instead of the PlanckBAO one.

External datasets – N_{eff}					
	Ω_M	h	$\ln(10^{10} A_s)$	n_s	N_{eff}
Fiducial values	0.3075	0.6774	3.064	0.9667	3.04
PlanckBAO	0.0088	0.015	0.038	0.0088	0.23
Euclid02	0.0031	0.0064	0.0136	0.0106	0.37
Euclidnl	0.0027	0.0041	0.0096	0.0056	0.23
CMB-S4	0.0012	0.00061	0.0094	0.0028	0.029

Table 10. Forecasted 1σ constraints on cosmological parameters and the N_{eff} considering each external dataset alone. These constraints are obtained using the Fisher matrices explained in Section 4. PlanckBAO represents Planck CMB combined with BOSS BAO measurements. We show the projected constraints for Euclid for two different k_{max} used and the constraints coming from CMB-S4 experiment.

Ext-HIRAX N_{eff}					
	Ω_M	h	$\ln(10^{10} A_s)$	n_s	N_{eff}
Fiducial values	0.3075	0.6774	3.064	0.9667	3.04
No wedge $k_{\text{max}} = 0.2h\text{Mpc}^{-1}$					
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0017	0.0036	0.013	0.0061	0.22
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0019	0.0037	0.023	0.0061	0.23
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0020	0.0037	0.041	0.0061	0.23
No wedge $k_{\text{max}} = k_{\text{nl}}(z)$					
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0010	0.0013	0.009	0.0014	0.075
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0011	0.0013	0.020	0.0015	0.077
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0011	0.0013	0.040	0.0015	0.077
Mid wedge $k_{\text{max}} = 0.2h\text{Mpc}^{-1}$					
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0018	0.0039	0.013	0.0067	0.24
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0021	0.0041	0.023	0.0067	0.25
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0022	0.0041	0.042	0.0067	0.25
Mid wedge $k_{\text{max}} = k_{\text{nl}}(z)$					
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0011	0.0014	0.010	0.0016	0.081
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0012	0.0014	0.021	0.0016	0.084
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0013	0.0014	0.040	0.0016	0.085
Wedge $k_{\text{max}} = 0.2h\text{Mpc}^{-1}$					
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0057	0.0117	0.03	0.0192	0.73
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0058	0.0118	0.04	0.0192	0.74
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0060	0.0119	0.06	0.0193	0.75
Wedge $k_{\text{max}} = k_{\text{nl}}(z)$					
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0027	0.0035	0.01	0.0040	0.21
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0031	0.0035	0.03	0.0041	0.21
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0035	0.0035	0.05	0.0042	0.22

Table 11. Fiducial values and 68% confidence intervals on cosmological parameters and the N_{eff} using Ext-HIRAX alone for various considerations. Going from top to bottom, we show the constraints considering different priors on both $b_{\text{HI}} \& \Omega_{\text{HI}}$, different k_{max} we use and different level of foreground wedge contamination.

highzFAST N_{eff}					
	Ω_M	h	$\ln(10^{10} A_s)$	n_s	N_{eff}
Fiducial values	0.3075	0.6774	3.064	0.9667	3.04
$k_{\text{max}} = 0.2h\text{Mpc}^{-1}$					
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0053	0.011	0.027	0.021	0.67
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0061	0.012	0.036	0.021	0.70
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0067	0.012	0.053	0.021	0.72
$k_{\text{max}} = k_{\text{nl}}(z)$					
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0038	0.0058	0.018	0.0078	0.34
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0046	0.0059	0.028	0.0079	0.36
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.0052	0.0060	0.047	0.0079	0.37

Table 12. Fiducial values and 68% confidence intervals on cosmological parameters and the N_{eff} using highzFAST alone for various considerations. We show the constraints considering different astrophysical priors on both $b_{\text{HI}} \& \Omega_{\text{HI}}$, as well as the dependence of the constraints on the k_{max} used.

Ext-HIRAX N_{eff}						
	$k_{\text{max}} = 0.2h\text{Mpc}^{-1}$			$k_{\text{max}} = k_{\text{nl}}(z)$		
Euclid Rec.+PlanckBAO	0.067			0.064		
+21cm	No wedge	Mid wedge	Wedge	No wedge	Mid wedge	Wedge
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.046	0.047	0.057	0.038	0.039	0.049
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.047	0.048	0.057	0.039	0.040	0.049
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.047	0.048	0.057	0.039	0.041	0.050
Euclid Rec.+CMB-S4	0.022			0.020		
+21cm	No wedge	Mid wedge	Wedge	No wedge	Mid wedge	Wedge
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.019	0.019	0.021	0.015	0.015	0.017
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.020	0.020	0.022	0.015	0.016	0.018
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.020	0.020	0.022	0.015	0.016	0.018

Table 13. Forecasted 1σ constraints on N_{eff} considering Ext-HIRAX instrument combined with CMB+galaxy probes. We show the constraints considering different priors on both $b_{\text{HI}} \& \Omega_{\text{HI}}$, different level of wedge contamination and different k_{max} used. In the top half of the table we show what are the gains of adding Ext-HIRAX measurements on top of combined Fisher matrix for Euclid galaxy survey and PlanckBAO. In the bottom half we do the same only using CMBS-4 Fisher matrix instead of the PlanckBAO one.

Ext-CHIME N_{eff}						
	$k_{\text{max}} = 0.2h\text{Mpc}^{-1}$			$k_{\text{max}} = k_{\text{nl}}(z)$		
Euclid Rec.+PlanckBAO	0.067			0.064		
+21cm	No wedge	Mid wedge	Wedge	No wedge	Mid wedge	Wedge
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.049	0.052	0.058	0.043	0.045	0.052
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.050	0.053	0.058	0.044	0.047	0.052
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.050	0.053	0.058	0.044	0.047	0.052
Euclid Rec.+CMB-S4	0.022			0.020		
+21cm	No wedge	Mid wedge	Wedge	No wedge	Mid wedge	Wedge
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.020	0.021	0.021	0.016	0.017	0.018
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.020	0.021	0.022	0.016	0.017	0.018
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.020	0.021	0.022	0.016	0.017	0.018

Table 14. Forecasted 1σ constraints on N_{eff} considering Ext-CHIME instrument combined with CMB+galaxy probes. We show the constraints considering different priors on both $b_{\text{HI}} \& \Omega_{\text{HI}}$, different level of wedge contamination and different k_{max} used. In the top half of the table we show what are the gains of adding Ext-CHIME measurements on top of combined Fisher matrix for Euclid galaxy survey and PlanckBAO. In the bottom half we do the same only using CMBS-4 Fisher matrix instead of the PlanckBAO one.

highzFAST N_{eff}		
	$k_{\text{max}} = 0.2h\text{Mpc}^{-1}$	$k_{\text{max}} = k_{\text{nl}}(z)$
Euclid Rec.+PlanckBAO	0.067	0.064
+21cm		
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.062	0.057
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.062	0.058
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.062	0.058
Euclid Rec.+CMB-S4	0.022	0.020
+21cm		
2% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.022	0.019
5% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.022	0.019
10% $b_{\text{HI}} \& \Omega_{\text{HI}}$	0.022	0.019

Table 15. Forecasted 1σ constraints on N_{eff} considering highzFAST instrument combined with CMB+galaxy probes. We show the constraints considering different priors on both b_{HI} & Ω_{HI} , different level of wedge contamination and different k_{max} used. In the top half of the table we show what are the gains of adding Ext-CHIME measurements on top of combined Fisher matrix for Euclid galaxy survey and PlanckBAO. In the bottom half we do the same only using CMBS-4 Fisher matrix instead of the PlanckBAO one.

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